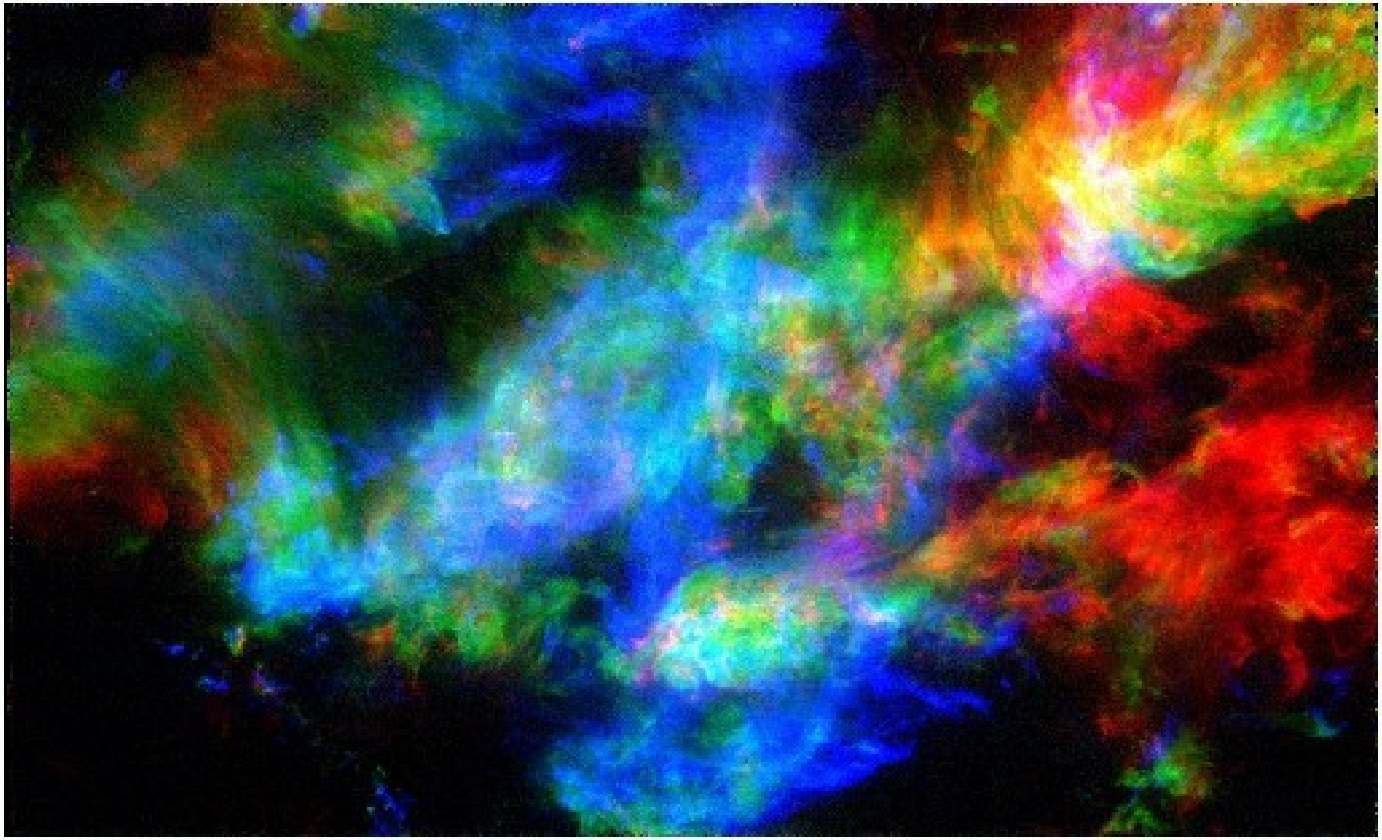
Molecules, mm & sub-mm (radio) astronomy



One of the last large projects completed with the FCRAO 14-m telescope was a map of 100 square degree region of the Taurus Molecular Cloud Complex in the emissions of the J=1-0 transitions of ^{12}CO and ^{13}CO. In addition to providing the distribution of molecular gas in this region, the high resolution spectra provide information on the cloud kinematics. The image below shows the ^{12}CO emission that is color coded, with red representing gas moving away, blue gas moving toward, and green gas moving at the average velocity of the cloud. Ron Snell

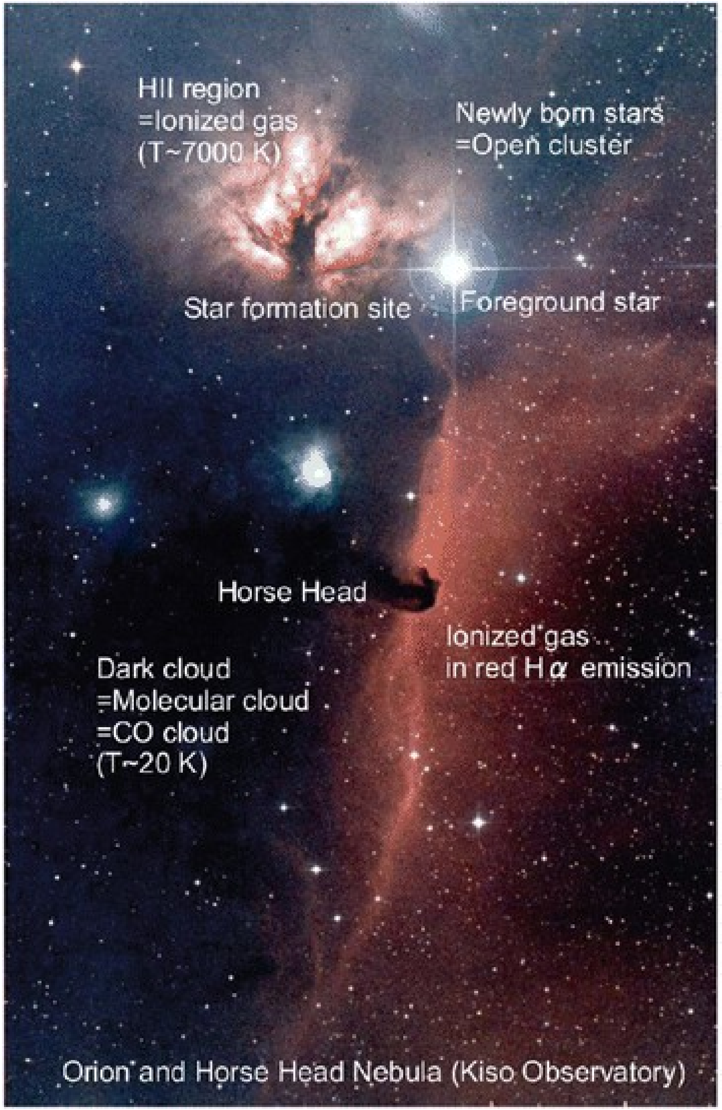
Molecules, mm & sub-mm (radio) astronomy

Further readings:

- Fanti & Fanti § 13.3 (only for generalities)
- Wilson, Rohlfs & Huttemeister "Tools of Radio Astronomy", chaps. 15-16
- Longair "High Energy Astrophysics", chap 12
- Heyer & Dame "Molecular Clouds in the Milky Way"
  2015, ARAA 53, 583-629 + a number of reviews cited at page 585
- Carilli & Walter "Cool Gas in High Redshift Galaxies"
  2013, ARAA 51, 105-161

Molecules, mm & sub-mm (radio) astronomy

Existence of dense molecular clouds is one of the first clues to understand the formation of stars



- First visually identified as dark nebulae by William and Caroline Herschel (1785),
- Photographic observations of Barnard (1919) and Wolf (1923) established as discrete, optically opaque interstellar **clouds**.
- In **1950, discovery of the HI line** (21 cm emission): correlation between dust absorption and HI emission (Lilley, 1955). BUT, observations towards the centers of dark nebulae detected either very weak or no HI emission
- **Bok** et al. (1955) suggested that within these nebulae the gas had to be in molecular form.
- Cold molecular component of the interstellar medium (ISM) was discovered in **1970** (Wilson, Jefferts & Penzias, ApJL 161, 43-44) via **CO observations**. "We have detected intense line radiation from the direction of the Orion Nebula at a frequency of 115,271.2 MHz"

Molecules, mm & sub-mm (radio) astronomy

- Immediately realized that **dark clouds were made of molecules**, mainly consisting of molecular hydrogen mixed with small amounts of **interstellar dust** and trace amounts of more complex molecular species.

- **Molecular clouds are the sites of all star formation** in the Milky Way

- Our understanding of molecular clouds comes from the Milky Way (*optimal spatial resolution, bad distance evaluation*) and from external galaxies (favorable viewing angles)

- For about 40 years most studies conducted with moderate (~15m) to large (30-45m) telescopes plus a few interferometers made of a handful of antennas. "Recently", ALMA provided resolution and sensitivity (collecting area!)

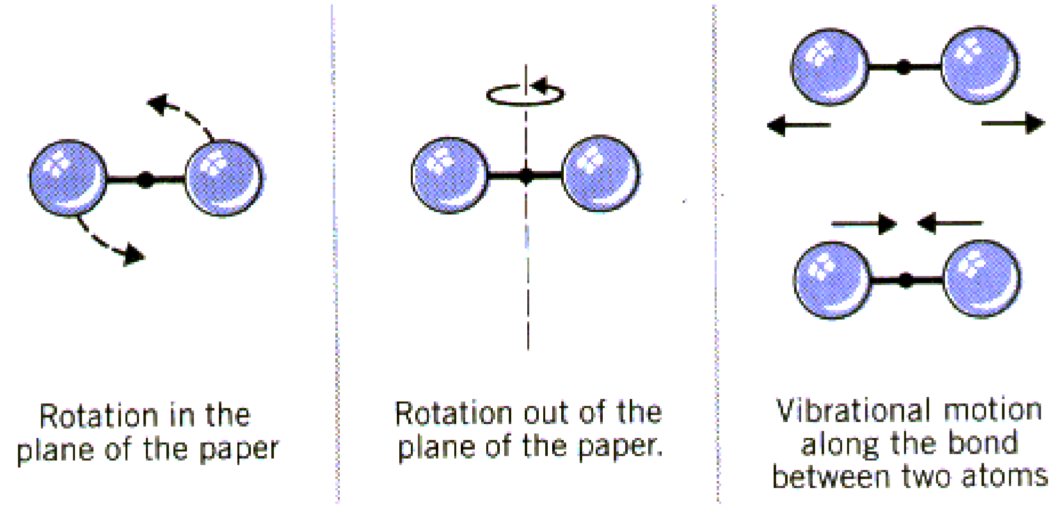

19 ALMA Antennas as of Oct. 2011

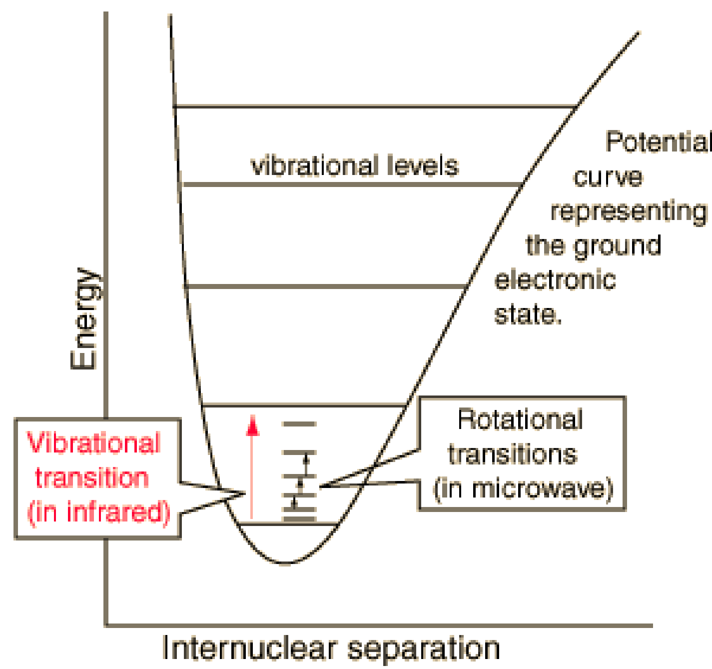
Molecules: what are we talking about?

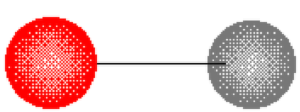
Molecules are ver difficult to deal with for quantum mechanics (positions and moments of all components, nuclei & electrons, must be included in the Schrodinger equation)

Transitions grouped into 3 sets

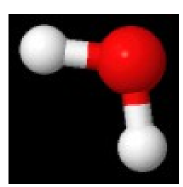
- *Electronic transitions (~ eV)*
- *Vibrational transitions (10 – 100 meV)*
  https://en.wikipedia.org/wiki/Molecular_vibration
- ***Rotational transitions** (~ meV)*







CO



H_2O

(Born-Oppenheimer approximation)

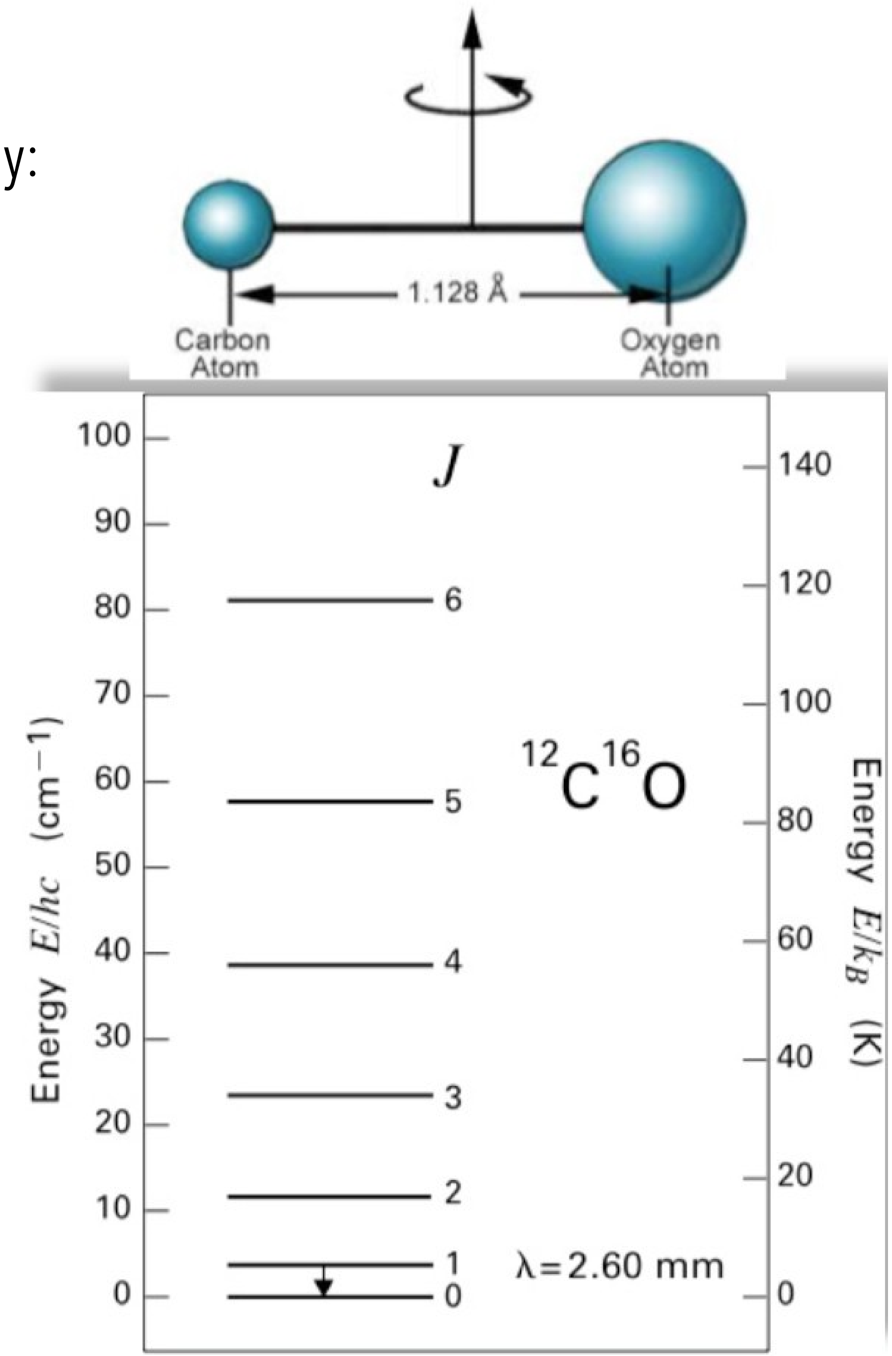
Quantum mechanics expression of the rotational energy:

$$E_{rot} = \frac{h}{2\pi I} J(J+1) \qquad \text{with} \quad \Delta J = \pm 1$$

The J=1 state is elevated above the ground by $4.8 \cdot 10^{-4}$ eV or, equivalently, 5.5 K!

Easy to excite in a quiescent cloud.

Within a molecular cloud, excitation of CO to the J=1 level occurs primarily through collisions with the ambient H_2.



Most common species: Molecules with permanent dipole momentum

Aka: non – homonuclear

- In the plane of rotation the dipole moment can be viewed as an antenna, oscillating as the molecule rotates.
- Classically, the acceleration of positive and negative charges gives rise to radiation whose frequency is that of rotation frequency.
- In the quantum mechanical model, the angular momentum is quantized, so that the radiation is emitted at discrete frequencies, when total angular momentum changes
- Allowed dipole radiative transitions will occur between different rotational states only if the molecule is polar.
- Homonuclear diatomic molecules (O_2 , H_2, N_2) cannot undergo allowed transitions, therefore they are very difficult to detect (needs asymmetric excitations of the electrons within the atoms).

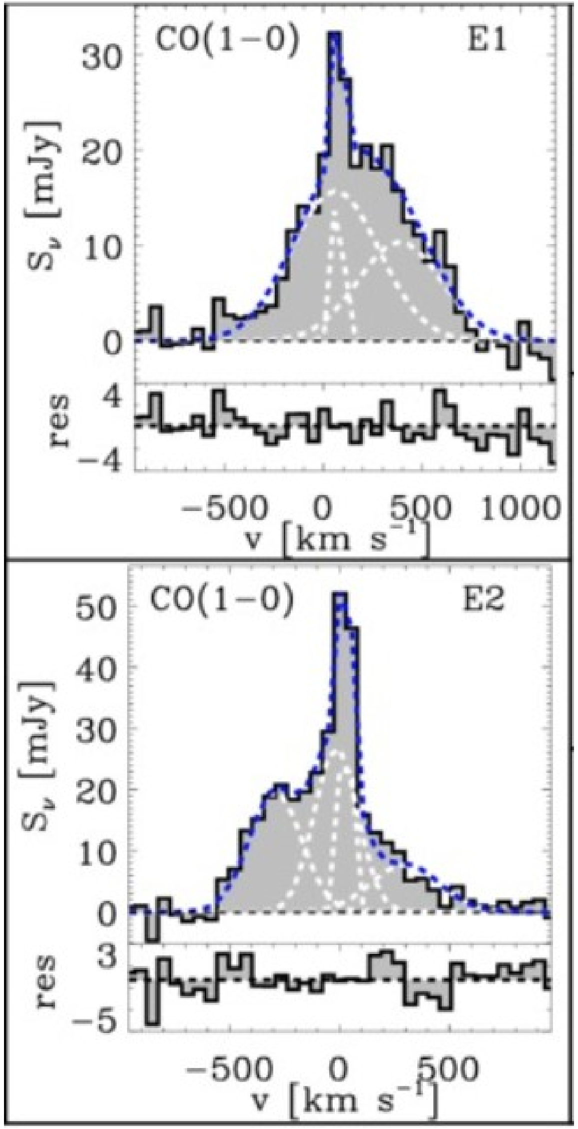
Molecular line measurement (emission/absorption): conversion of line intensity, integrated over linewidth, into column density.

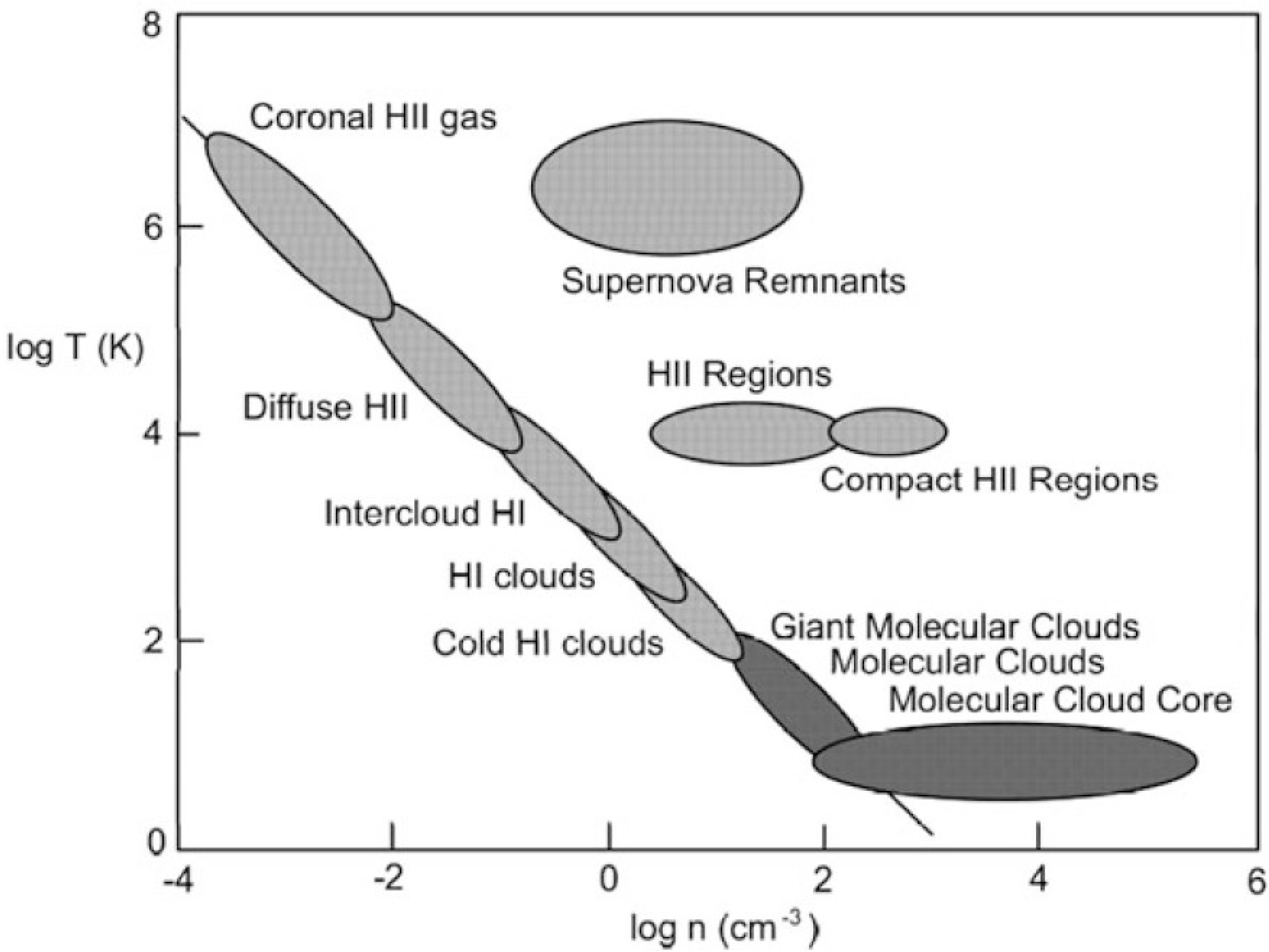
$$M_{CO} = \text{Factor}\, D^2 \int_{line} I(\nu)\, d\nu \qquad \text{units}$$

Aims:

1. Determine the complete chemical and isotopic content of a molecular cloud.

2. Relate the molecular clouds parameters to star formation
   2.1 Interaction with very young stellar objects
   2.2 Search for prebiotic molecules.




8
6
4
2
0
log T (K)
Coronal HII gas
Supernova Remnants
HII Regions
Compact HII Regions
Diffuse HII
Intercloud HI
HI clouds
Cold HI clouds
Giant Molecular Clouds
Molecular Clouds
Molecular Cloud Core
-4
-2
0
2
4
6
log n (cm^{-3})

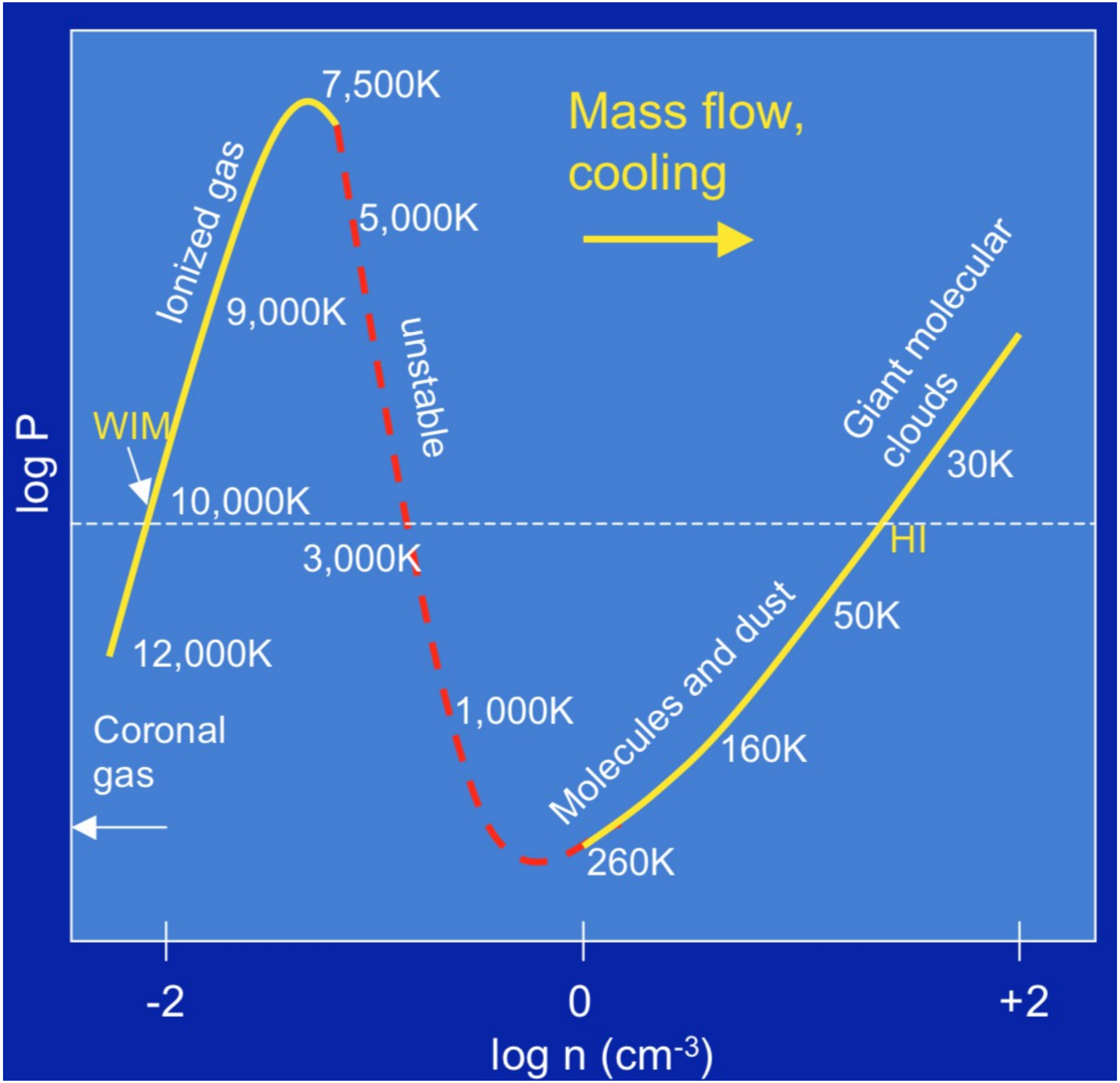
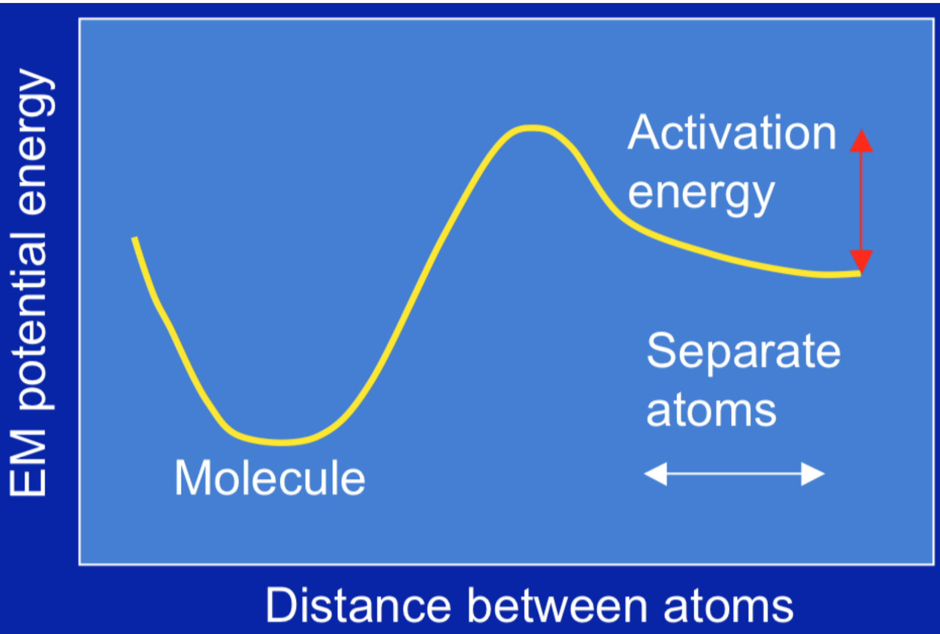
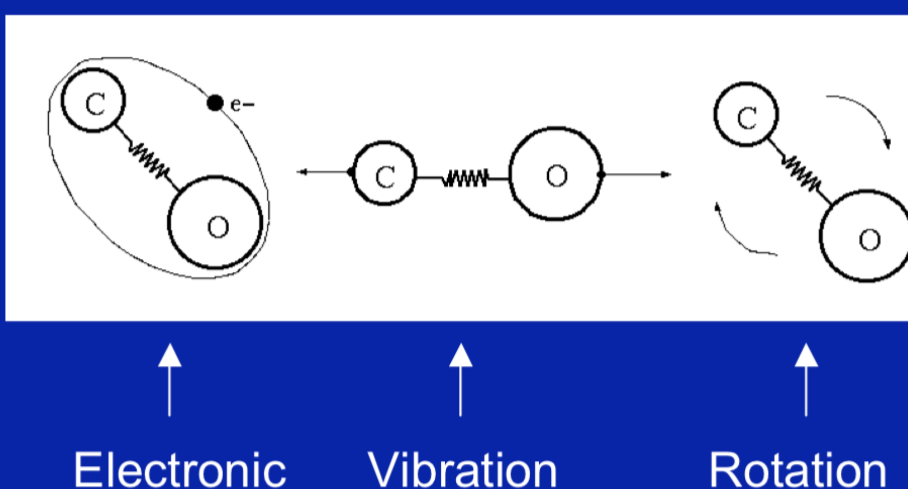

7,500K
Mass flow,
cooling
Ionized gas
5,000K
9,000K
unstable
Giant molecular
clouds
WIM
30K
log P
10,000K
HI
3,000K
50K
Molecules and dust
12,000K
1,000K
160K
Coronal
gas
260K
-2
0
+2
log n (cm^{-3})

Table 2.1 ISM and radiations

T(K), E	Object	Line	Continuum
2.7 K	CBR		mm,submm
10–50 K	Mol. cloud H_2 gas Dust	Mol. lines mm,submm	FIR
100–1000 K	HI clouds Diffuse HI Dust	21-cm	IR
10^4 K	HII regions	Recom.lines	Free-free
$10^{5\sim7}$ K	SNR Diffuse HII Halo		Synch. X-rays
High-energy (keV-GeV)	Mag. fi, CR SNR, Pulsars Gal. Center AGN, Jets Radio gal.		Synch.Polari.
γ rays (GeV-TeV)	CR		γ rays

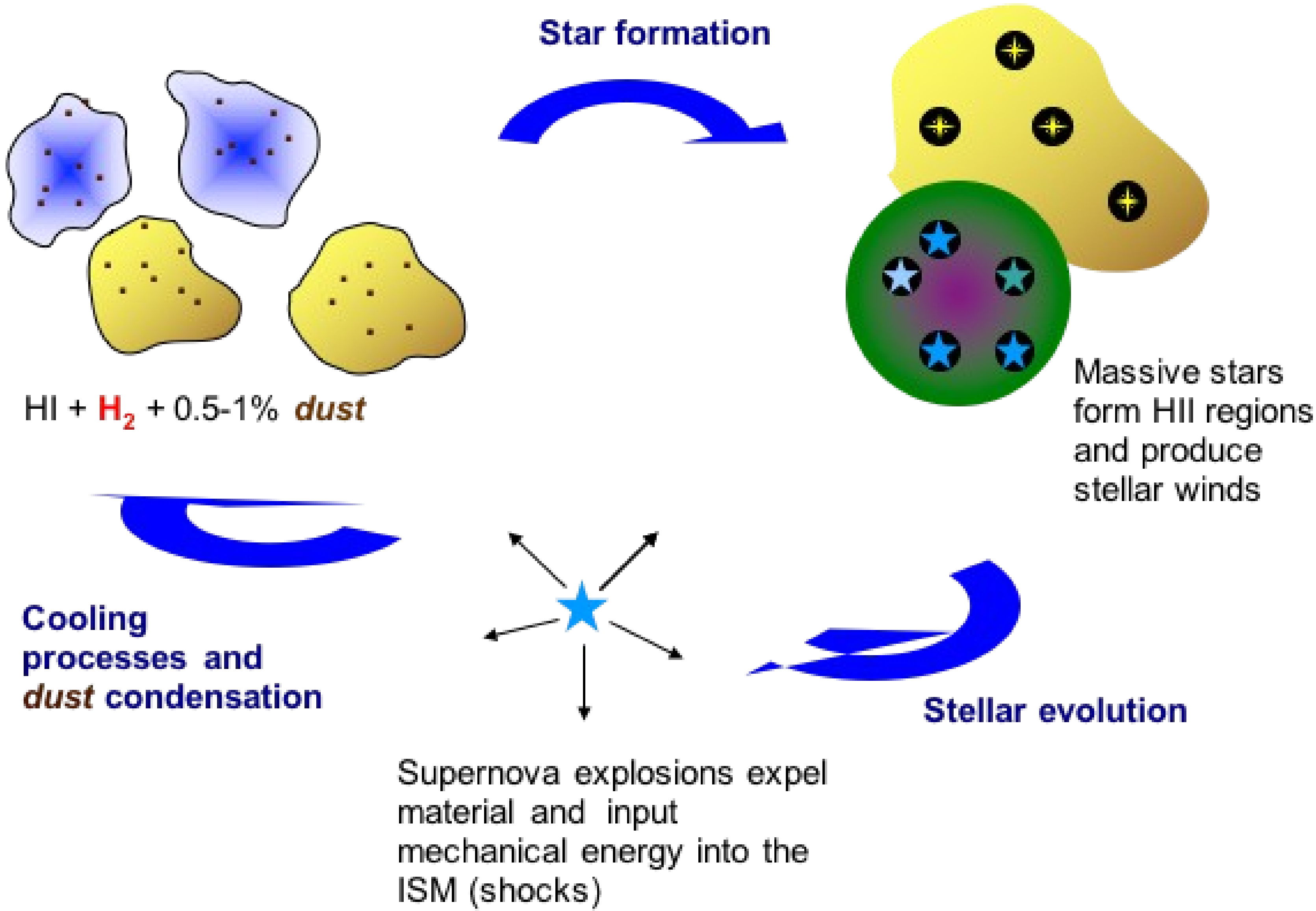




- Chemical bonding involves sharing electrons, therefore it uses electromagnetic, not nuclear forces
- Stable molecule has to have lower EM potential energy than the sum of separate atoms that make it up
- Electrostatic repulsion of electrons form a barrier called the activation energy
- Excitation and energy release by a molecule occurs through one of the three mechanisms: electronic, vibration, and rotation with rotation being important in the gas phase and vibration dominating in the solid state environment (ices on top of dust grains)

- Gas-phase, ion-molecule: cosmic rays ionize H_2, make H_2^+, H^+ (also He^+), which react with H_2 and CO (most abundant) and create simple neutral molecules and metal ions through successive reactions [$A^+ + B \rightarrow C^+ + D$]. Important in both diffuse and dense clouds
- Shock-induced: gas heating caused by shock wave overcomes activation barriers, enables neutral-neutral reactions [$A + B \rightarrow C + D$] that are impossible in cold clouds, may cause grain disruption. Important in dense, hot star-forming regions
- Circumstellar: "freeze-out" chemistry in circumstellar envelopes, gas-phase molecules stick to grains, which develop icy mantles. Reactions proceed on grain surfaces (hydrogenation of O, C, N, formation of H_2 most common, because of H mobility)
- Dust grain surfaces: shield molecules from UV radiation field, produce H_2 through catalysis: $H + H + \text{grain} \rightarrow H_2 + \text{grain}$ (reaction is exothermic, energy is used to detach H_2 from surface). H_2 drives much of gas-phase chemistry

The interstellar medium energy cycle



- Stars form from dense , cool H_2
- The ISM plays an important role in the energy exchange and provide raw material
- Star formation occurs in dense dusty clouds so need some tracer that penetrates the dust ("extinction free" trace)
- Spectral observations also trace kinematics, so can be used to probe dynamical mass, gas motion, and gravitational torques
- Serves as fuel for both starburst and AGN activity
- Significant mass in galaxy nuclei
  - H_2 is a "silent" molecule –need tracer species (CO)

Observability of H_2 emission

- Molecular hydrogen is a symmetric molecule

– The electric dipole moment is zero

– Rotational transitions are forbidden, even though electronic transitions are allowed

- Quadrupole transitions are possible

– They are very weak and characterized by energies much higher than the typical rotational energies

– Can only be observed in relatively warm regions

Thanks to the abundance of H_2 some of these transitions have been observed in the mid IR

- Lack of H_2 emissions $\Rightarrow$ hard to make maps of the distribution of molecular hydrogen

– UV absorptions require (rare) bright background sources and are not suited to make maps, especially in regions rich of dust and molecules, where the background sources are obscured

MOLECULAR LINE LIST: H_2, HD, H_2O, OH, CO

Compiled by E.F. van Dishoeck July 19, 1994

Species	Transition	Lambda (mu)	E_l (cm-1)	Reference
H_2	(1,0) Q(1)	2.4065914	118.49	Bragg et al. 1982
H_2	(2,1) Q(1)	2.5510	118.49	Black & van Dishoeck 1987
H_2	(2,1) Q(3)	2.5698	705.52	Black & van Dishoeck 1987
H_2	(2,1) Q(5)	2.6040	1740.24	Black & van Dishoeck 1987
H_2	(1,0) O(2)	2.6269	354.37	Black & van Dishoeck 1987
H_2	(1,0) O(3)	2.8025	705.52	Black & van Dishoeck 1987
H_2	(1,0) O(5)	3.2350	1740.24	Black & van Dishoeck 1987
H_2	(0,0) S(13)	3.8468	9580.75	Jennings et al. 1985 constants
H_2	(0,0) S(11)	4.1813	7148.71	Jennings et al. 1985 constants
H_2	(0,0) S(10)	4.4099	6039.15	Jennings et al. 1985 constants
H_2	(0,0) S(9)	4.69461	5005.73	Jennings et al. 1987
H_2	(0,0) S(8)	5.05303	4053.51	Jennings et al. 1987
H_2	(0,0) S(7)	5.51116	3188.05	Jennings et al. 1987
H_2	(0,0) S(6)	6.10856	2415.08	Jennings et al. 1987
H_2	(0,0) S(5)	6.90952	1740.24	Jennings et al. 1987
H_2	(0,0) S(4)	8.02505	1168.81	Jennings et al. 1987
H_2	(0,0) S(3)	9.66491	705.52	Jennings et al. 1987
HD	(0,0) R(10)	11.57346	4627.07	Ulivi et al. 1991
H_2	(0,0) S(2)	12.27861	354.37	Jennings et al. 1987
HD	(0,0) R(9)	12.47181	3825.27	Ulivi et al. 1991
HD	(0,0) R(8)	13.59265	3089.58	Ulivi et al. 1991
HD	(0,0) R(7)	15.25104	2423.88	Ulivi et al. 1991
HD	(0,0) R(6)	16.89381	1831.95	Ulivi et al. 1991
H_2	**(0,0) S(1)**	**17.03483**	118.49	Jennings et al. 1987 <- lowest energy, corresponding to T_excitation ~ 850 K
p-H_2O	5_51-4_04	19.2300	222.05	HITRAN
HD	(0,0) R(5)	19.43100	1317.31	Ulivi et al. 1991
o-H_2O	5_50-4_23	22.6391	300.36	HITRAN
HD	(0,0) R(4)	23.03376	883.16	Ulivi et al. 1991
OH 1/2-3/2	9/2-7/2	24.614	201.93	Offer & van Dishoeck 1992
OH 1/2-3/2	9/2-7/2	24.642	202.38	Offer & van Dishoeck 1992
o-H_2O	5_41-4_14	25.9402	224.84	HITRAN
H_2	(0,0) S(0)	28.21883	0.00	Jennings et al. 1987
HD	(0,0) R(3)	28.50197	532.31	Ulivi et al. 1991
p-H2O	4_40-3_13	28.9138	142.28	HITRAN
OH 1/2-3/2	7/2-5/2	28.939	83.72	Offer & van Dishoeck 1992
OH 1/2-3/2	7/2-5/2	28.940	83.92	Offer & van Dishoeck 1992

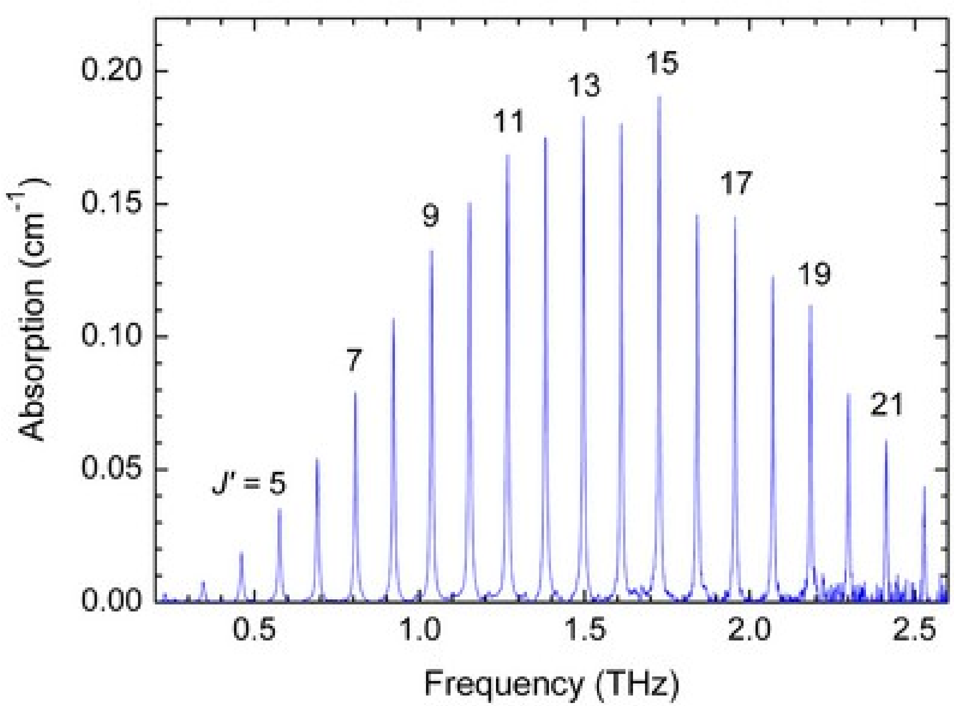
Estimating the exctitation temperature, in K, of a given transition (using MKs units)

$$T = \frac{h\nu}{k} = \frac{hc}{\lambda k} = \frac{6.626\cdot 10^{-34}\,\mathrm{kg\,m^2\,s^{-1}}}{1.38\cdot 10^{-23}\,\mathrm{kg\,m^2\,s^{-2}\,K^{-1}}}\,\nu =$$

$$= 4.80\cdot 10^{-11}\,\nu = 4.80\cdot 10^{-11}\,\frac{c}{\lambda} \quad [K]$$

Example of molecular transitions (model) for the CO

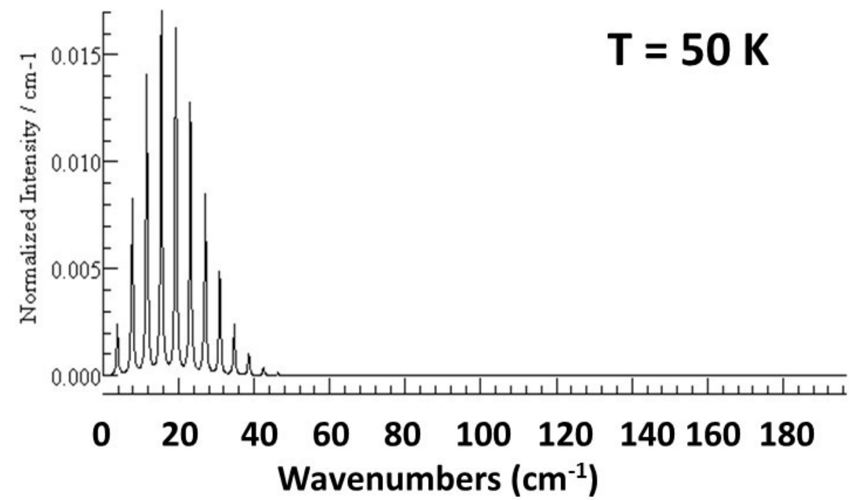
A rigid rotator is the model assumed



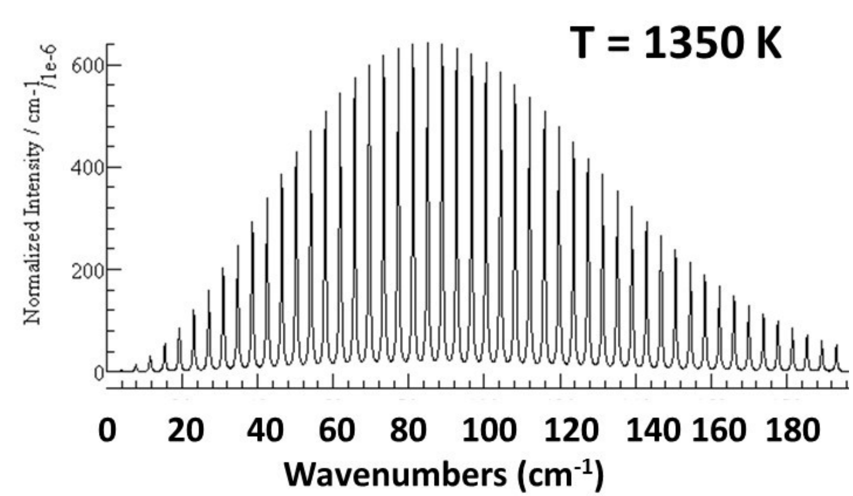
Laboratory CO spectrum in a 2 bar cell

CO Pure Rotational Spectrum

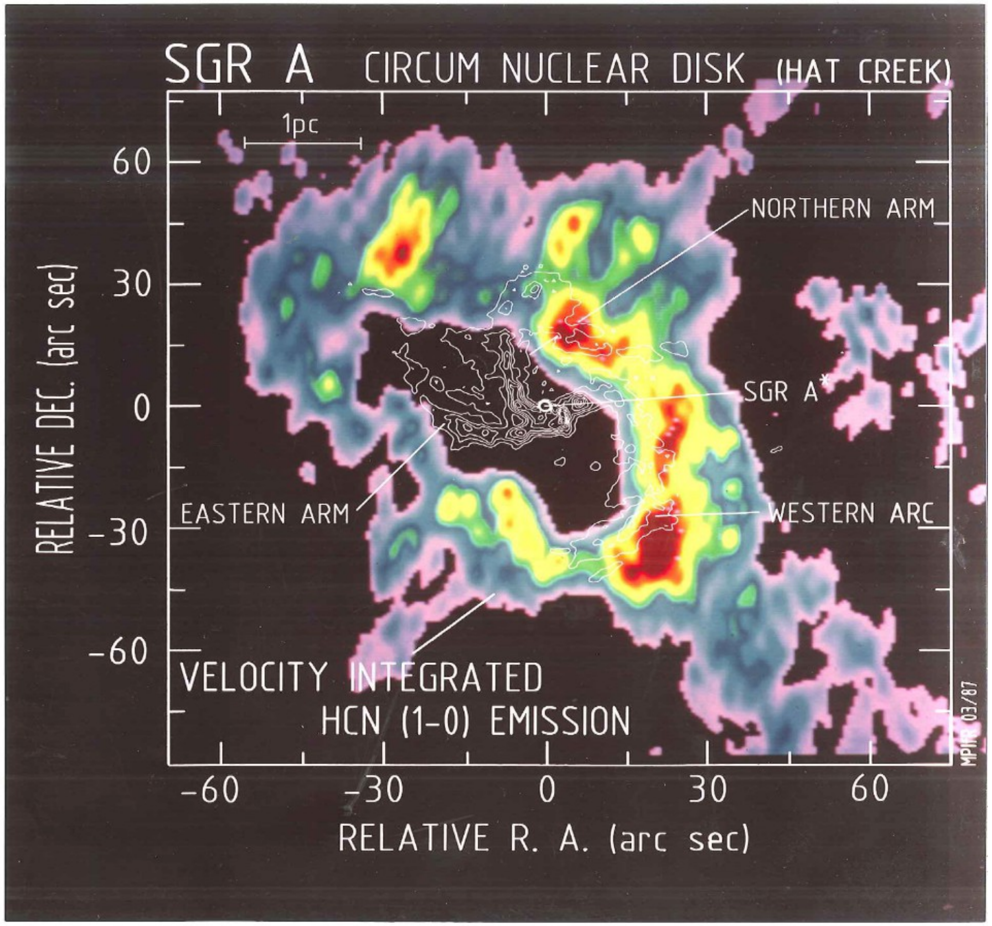
Temperature dependence due to the population of rotational energy levels.



The maximum in J scales with kT.



Complex molecules (3+ elements) have more and more complicated energy levels, depending on geometry.
There is coupling with the electric dipole/quadripole and magnetic dipole of the e- and nuclei and all this produces hyperfine transitions, splitting single levels in multiple line profiles with tiny separation in energy.



The energy separation (as small as a few MHz) depends on the position of a given nucleus in the molecule.

e.g. HCN and HNC have different level separation

All dipole transitions are allowed and will occur only in case the molecule is polar

Homonuclear molecules (O_2, H_2, N_2, C_2,...) difficult to detect and have forbidden transitions:

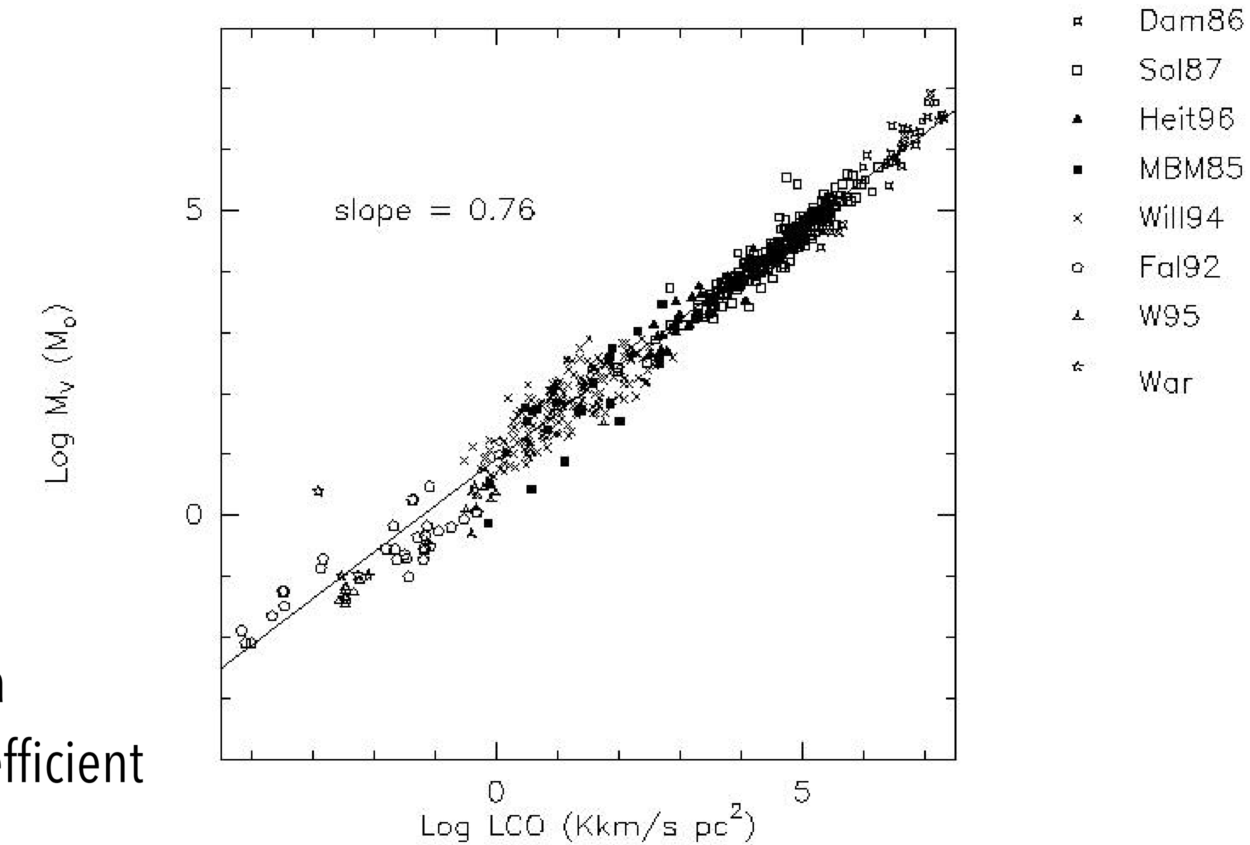
H_2 has large spacings of the rotational levels: first quadrupole line ~ 500 K above ground level (inefficient radiation from most of the cold ISM (where T < 200 K)

Detected/evaluated via tracers

- Extinction + thermal continuum emission from dust grains
- Gamma rays from CR collisions
- Molecular line spectroscopy (^{12}CO, ^{13}CO, $C^{18}O$)

CO has a small dipole moment and a correspondingly small Einstein A coefficient

CO is the best tracer of H2



Virial mass .vs. CO luminosity

Critical density:

$$\frac{A_{ul}}{n_p Q_{ul}} = 1 \quad \rightarrow \quad \text{namely} \quad \rightarrow \quad n_p^{cr} = \frac{A_{ul}}{Q_{ul}}$$

It is different for each transition! CO J1 → 0, has a critical density of ~ 2000 cm^{-3}.

Dense regions ensure a consistent population of the excited state (kinetic temperature)

CO forms in regions where visual extinction is > 1-3 Mag, dense enough that the line is optically thick: i.e. Emitted photons absorbed by another molecule

^{12}CO complemented with ^{13}CO (25-100 less abundant) ....
moderate optical depth, and then more appropriate to determine H_2 in dense regions

CO and H_2 are linked since formed and destroyed by similar physical processes:

Formation: surface of dust grains

Dissociation: H_2

1. incident UV photon excite molecule to Lyman ($\lambda < 1108$ Å) or Werner ($\lambda < 1008$ Å) bands

2. ~15% radiative de excitation into unbound within the electronic ground level that dissociated the molecule

Penetration of UV photons limited by dust attenuation + opacity provided by the outer H_2 layer ($\tau \sim = \sim \mu$ l, then deeper in the cloud higher τ, all H_2 molecular)

Competition between radiative and collisional de-excitation quantified by ratio C_{ul}/A_{ul}

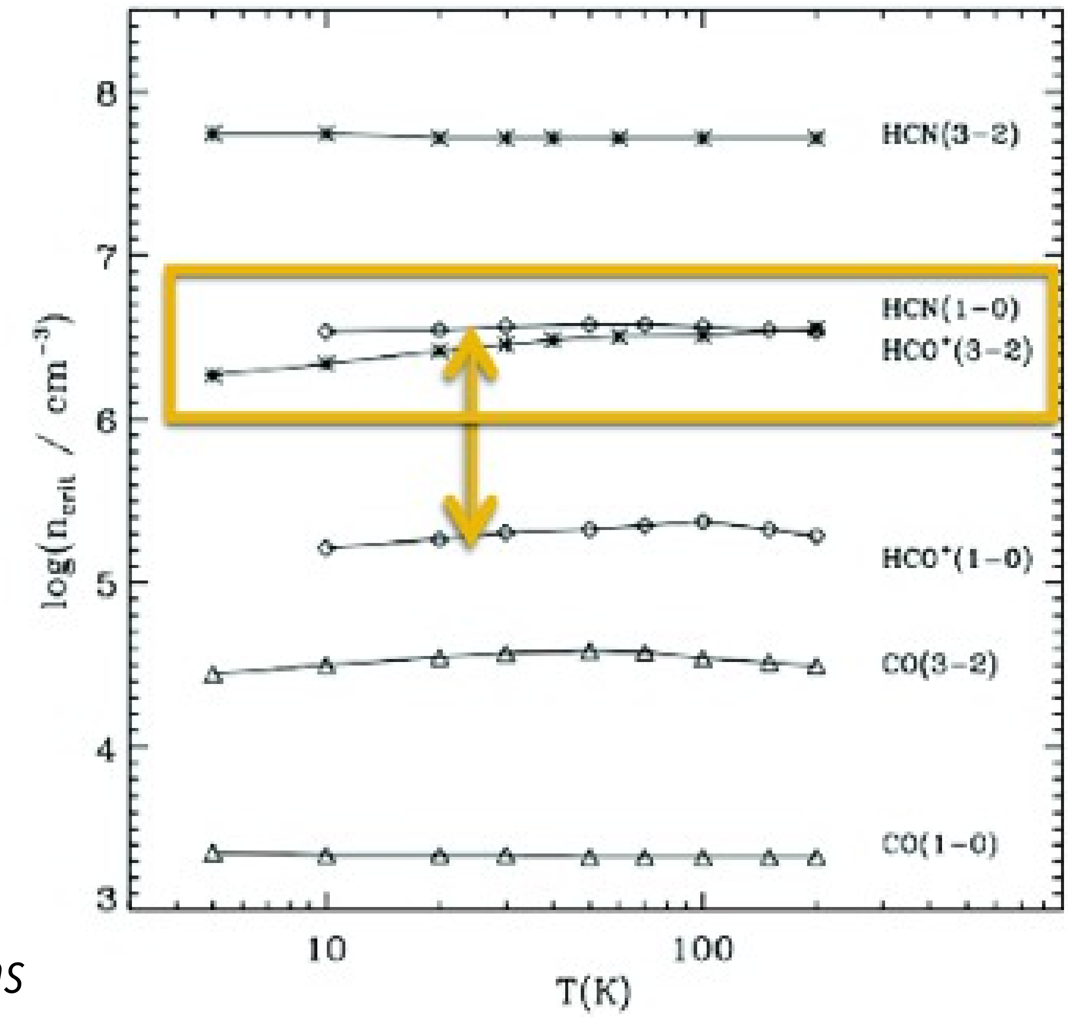
Critical density n_{crit} for which

$$A_{ul}/C_{ul} = 1 \sim J^3 m_{el}^2 / n(H_2) T^{1/2}$$

m_{el} = electric dipole moment,
J = angular momentum quantum number

*Highest n_{crit} molecules with largest m_{el} trace **hottest/densest** molecular gas.*

*On the same molecule, highest-order transitions trace **hottest/densest** molecular gas*



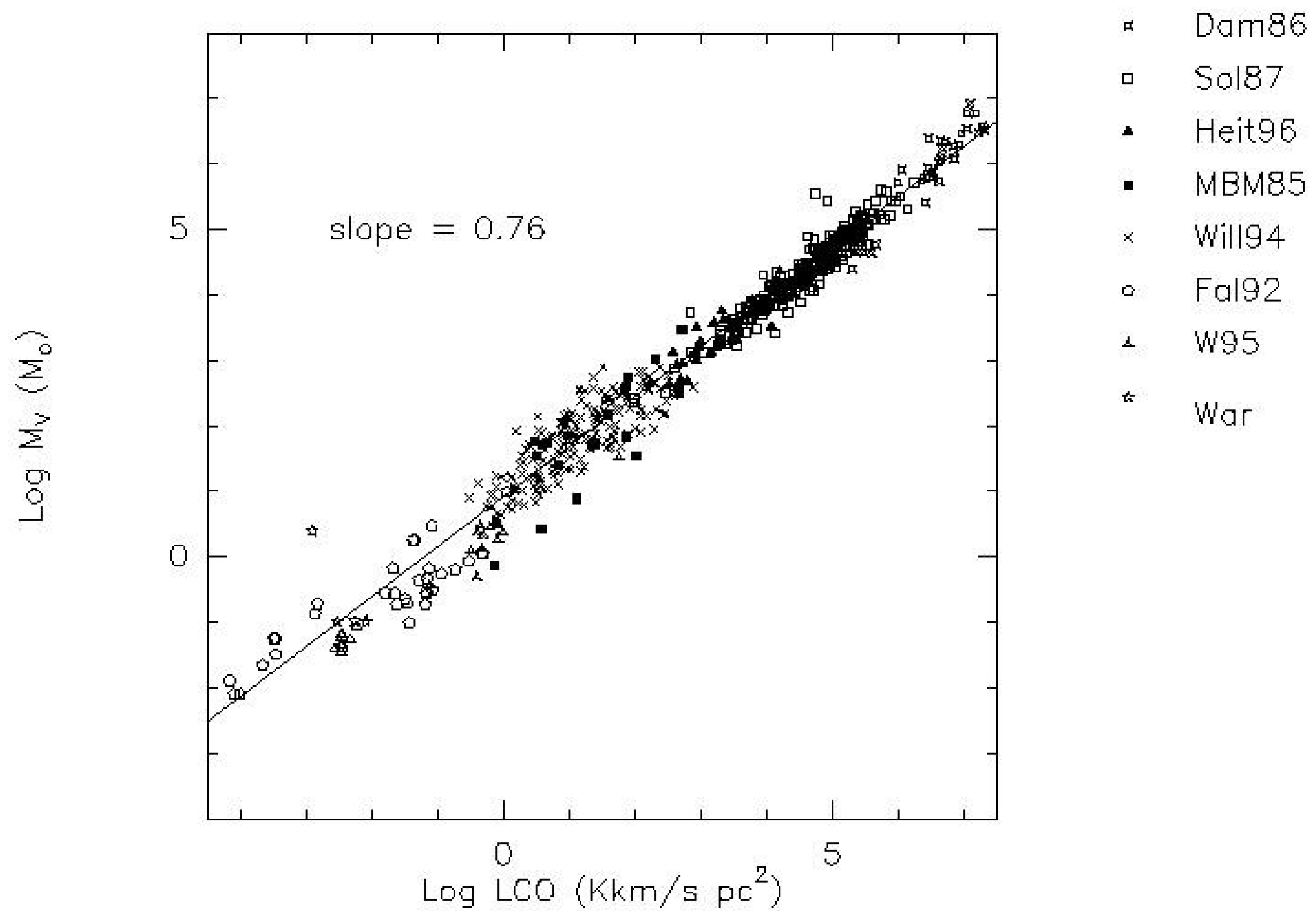
➲➲➲ *CO emission is optically thick (e.g., Wilson+ 1974), hence traces surface area Not suitable for estimating total gas amount. Need optically thin emission (all photons detected)*

Critical density & temperature

Species	Transition	Excitation potential (K)	ć (Õm)	č (GHz)	Einstein A (s^{-1})	n_{crit} (cm^{-3})
CO	J = 1–0	5.5	2601	115.27	7.2×10^{-8}	2.1×10^{3}
	J = 2–1	16.6	1300	230.54	6.9×10^{-7}	1.1×10^{4}
	J = 3–2	33.2	867	345.80	2.5 × 10−6	3.6×10^{4}
	J = 4–3	55.3	650.3	461.04	6.1 × 10−6	8.7×10^{4}
	J = 5–4	83.0	520.2	576.27	1.2 × 10−5	1.7×10^{5}
	J = 6–5	116.2	433.6	691.47	2.1 × 10−5	2.9×10^{5}
	J = 7–6	154.9	371.7	806.65	3.4 × 10−5	4.5×10^{5}
	J = 8–7	199.1	325.2	921.80	5.1 × 10−5	6.4×10^{5}
	J = 9–8	248.9	289.1	1036.9□ □	7.3 × 10−5	8.7×10^{5}
	J = 10–9	304.2	260.2	1152.0□ □	1.0 × 10−4	1.1×10^{6}
HCN	J = 1–0	4.25	3383	□88.63	2.4 × 10−5	2.6×10^{6}
	J = 2–1	12.76	1691	177.26	2.3 × 10−4	1.8 × 107
	J = 3–2	25.52	1128	265.89	8.4 × 10−4	6.8 × 107
	J = 4–3	42.53	845.7	354.51	2.1 × 10−3	1.8 × 108
	J = 5–4	63.80	676.5	443.12	4.1 × 10−3	3.8 × 108
	J = 6–5	89.32	563.8	531.72	7.2 × 10−3	7.1 × 108
	J = 7–6	119.09	483.3	620.30	1.2 × 10−2	1.2 × 109

Detected/evaluated via tracers

- Extinction + thermal continuum emission from dust grains
- Gamma rays from CR collisions
- Molecular line spectroscopy (^{12}CO, ^{13}CO,$C^{18}O$)



Virial mass .vs. CO luminosity

How to find the H_2 abundance: the X_{CO} factor:

$$X_{CO} = 1.6 \cdot 10^{19} \frac{T_{ex}}{30K} e^{\frac{5.5K}{T_{ex}} - 0.184} \quad cm^{-2} (K\,km\,s^{-1})^{-1}$$

Method	X_{CO} / 10^{20} cm^{-2}(K km s^{-1})$^{-1}$	References
Virial	2.1	Solomon et al. (1987)
	2.8	Scoville et al. (1987)
Isotopologues	1.8	Goldsmith et al. (2008)
Extinction	1.8	Frerking, Langer & Wilson (1982)
	2.9-4.2	Lombardi, Alves & Lada (2006)
	0.9-3.0	Pineda, Caselli & Goodman (2008)
	2.1	Pineda et al. (2010)
	1.7-2.3	Paradis et al. (2012)
Dust Emission	1.8	Dame, Hartmann & Thaddeus (2001)
	2.5	Planck Collaboration et al. (2011a)
γ-rays	1.9	Strong & Mattox (1996)
	1.7	Grenier, Casandjian & Terrier (2005)
	0.9-1.9 *	Abdo et al. (2010d)
	1.9-2.1 *	Ackermann et al. (2011 , 2012d)
	0.7-1.0 *	Ackermann et al. (2012b, 2012c)

H_2 column density

$$N_{H_2} = X_{CO} W_{CO} = X_{CO} \left(\int T_B(v)\, dv \; \mathrm{K\,km\,s^{-1}} \right) \quad \mathrm{cm^{-2}}$$

Fact:

➲➲ Stars form in high density regions

High-density molecular tracers: HCN(1-0)

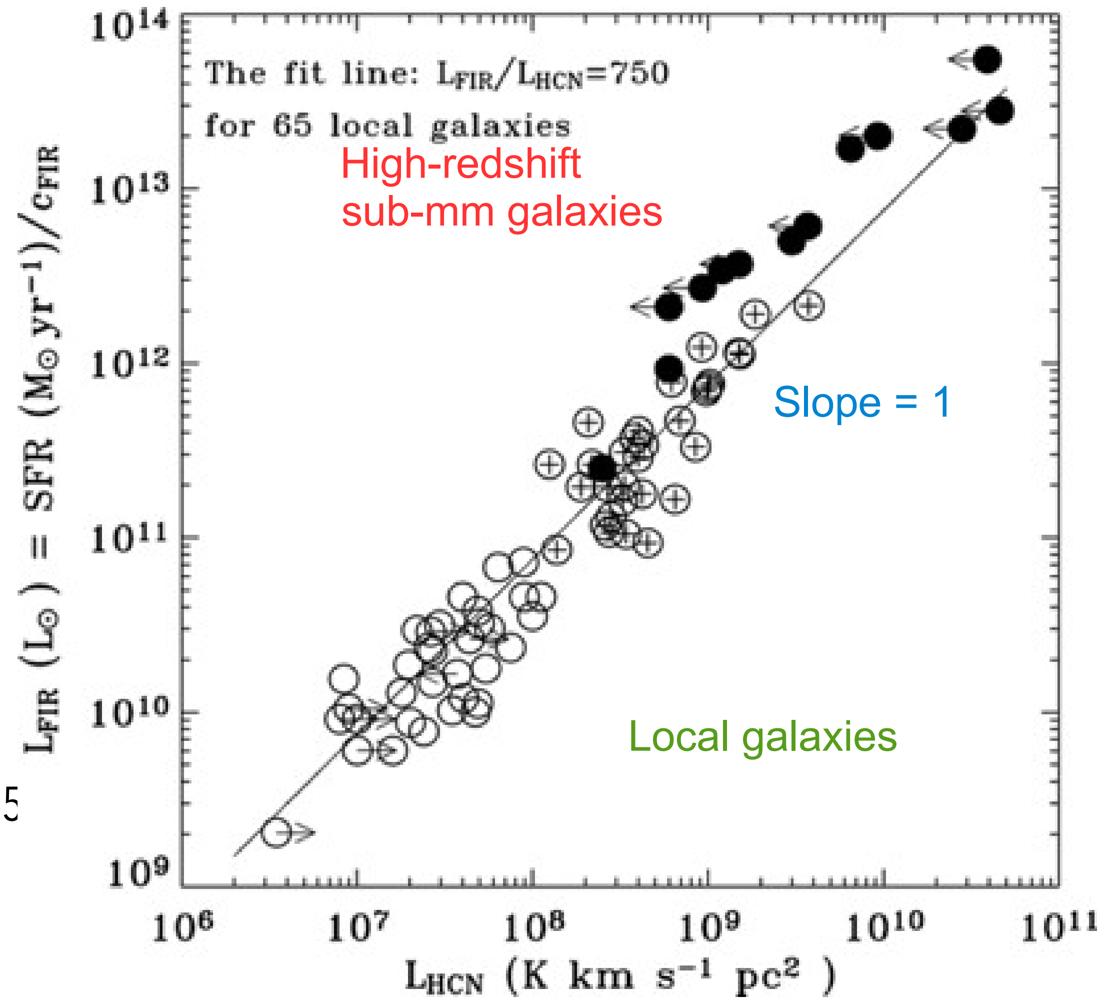
Correlation:

IR luminosity (~ SFR) and

HCN J=1-0 luminosity

(~mass of gas above critical density)

(Gao & Solomon 2004, Gao+ 2007, see also Wu+ 2005



Linear (rather than super-linear as for CO);

dense gas more closely correlated with star formation.

Fact: $A_{ul}/C_{ul} = 1 \sim J^3 m_{el}^2 / n(H_2) T^{1/2}$

Einstein A_{ul} coefficient ~ $J^3\ m_{el}^2$

i.e. : dense regions are transparent to high m_{el}^2 & high J transitions

➲ Best dense gas tracers :

CS (m_{el} = 1.958 Debye),

HCO^+ (m_{el} = 3.93 Debye)

HCN (m_{el} = 2.985 Debye)

NH3 (m_{el} = 1.48 Debye)

CO (m_{el} = 0.110 Debye) is less favorable, and it is a good tracer of moderately dense regions

➲ Which molecule is more convenient is determined by relative abundances (and observational requirements as well!)

Which molecule is more convenient is determined by relative abundances
(and observational requirements as well!)

CS (Carbon Monosulfide)

transitions in (J1-0 @ 48.99 GHz) mm & submm bands, level separation a few tens of GHz

HCO^+ ()

HCN (Hydrogen Cyanide)

transitions in (J1-0 @ 88.63 GHz) mm & submm bands, level separation a few tens of GHz

HNC (Hydrogen Isocyanide)

transitions in (J1-0 @ 90.66 GHz) mm & submm bands, level separation a few tens of GHz

NH_3 (para & otho ammonia)

transitions in radio (many between 23 & 25 GHz, wide range of temperatures) excellent to study protostellar cores !

Parameters for transitions of the most common interstellar molecules in table 16.1, Tool of Radio Astronomy

Molecular tracers of dense regions

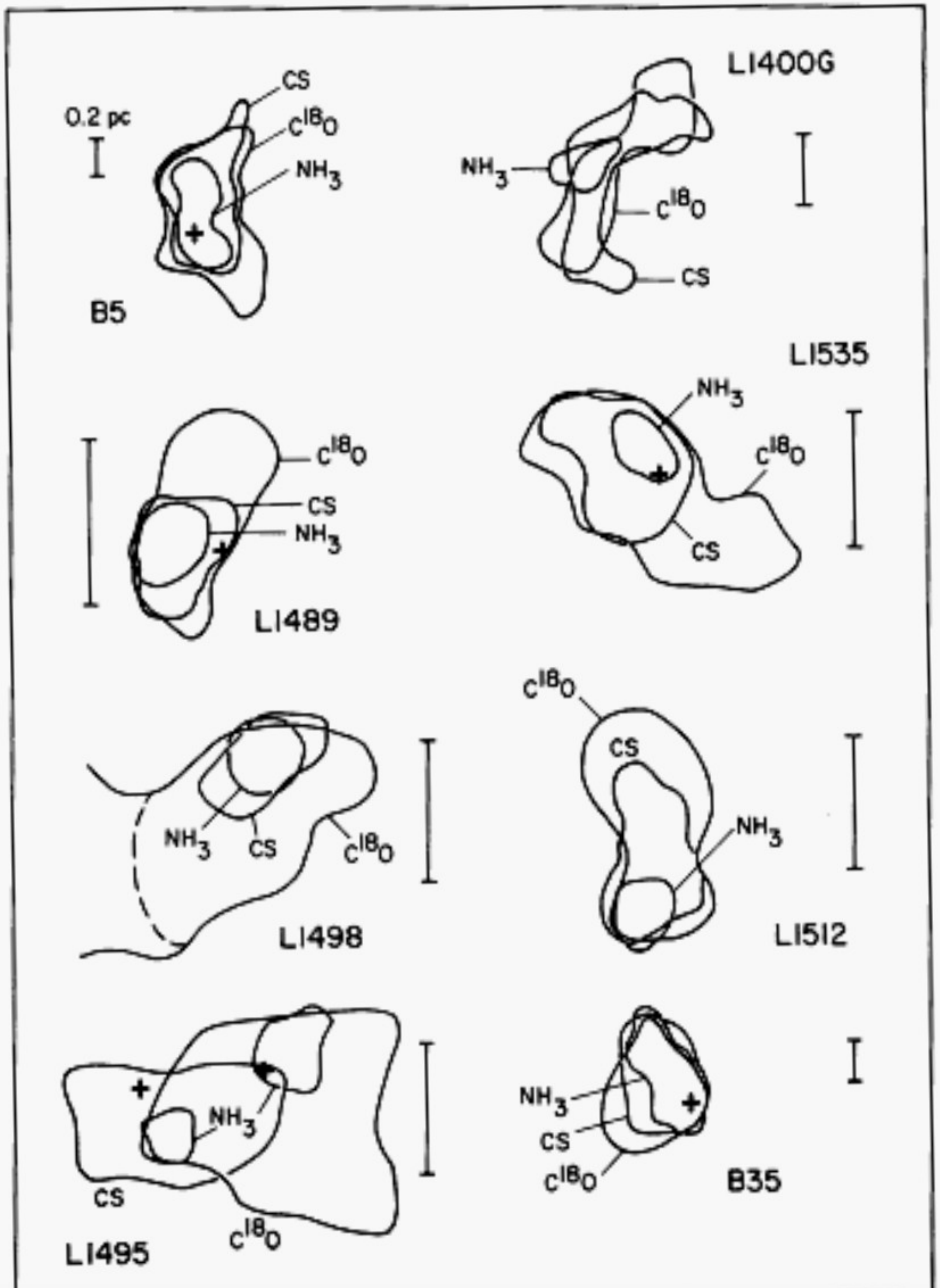


Fig. 1a

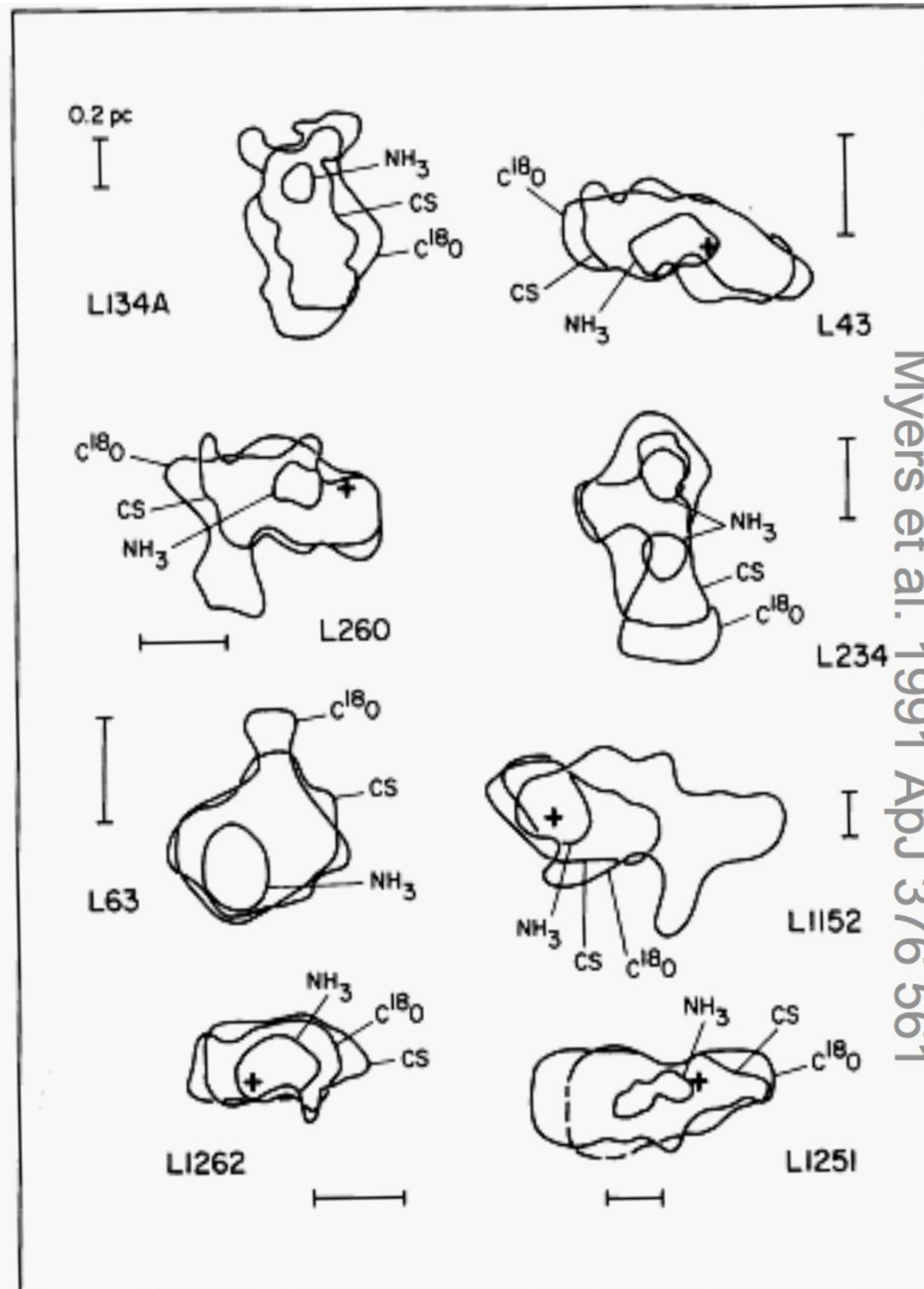


Fig. 1b

Myers et al. 1991 ApJ 376 561

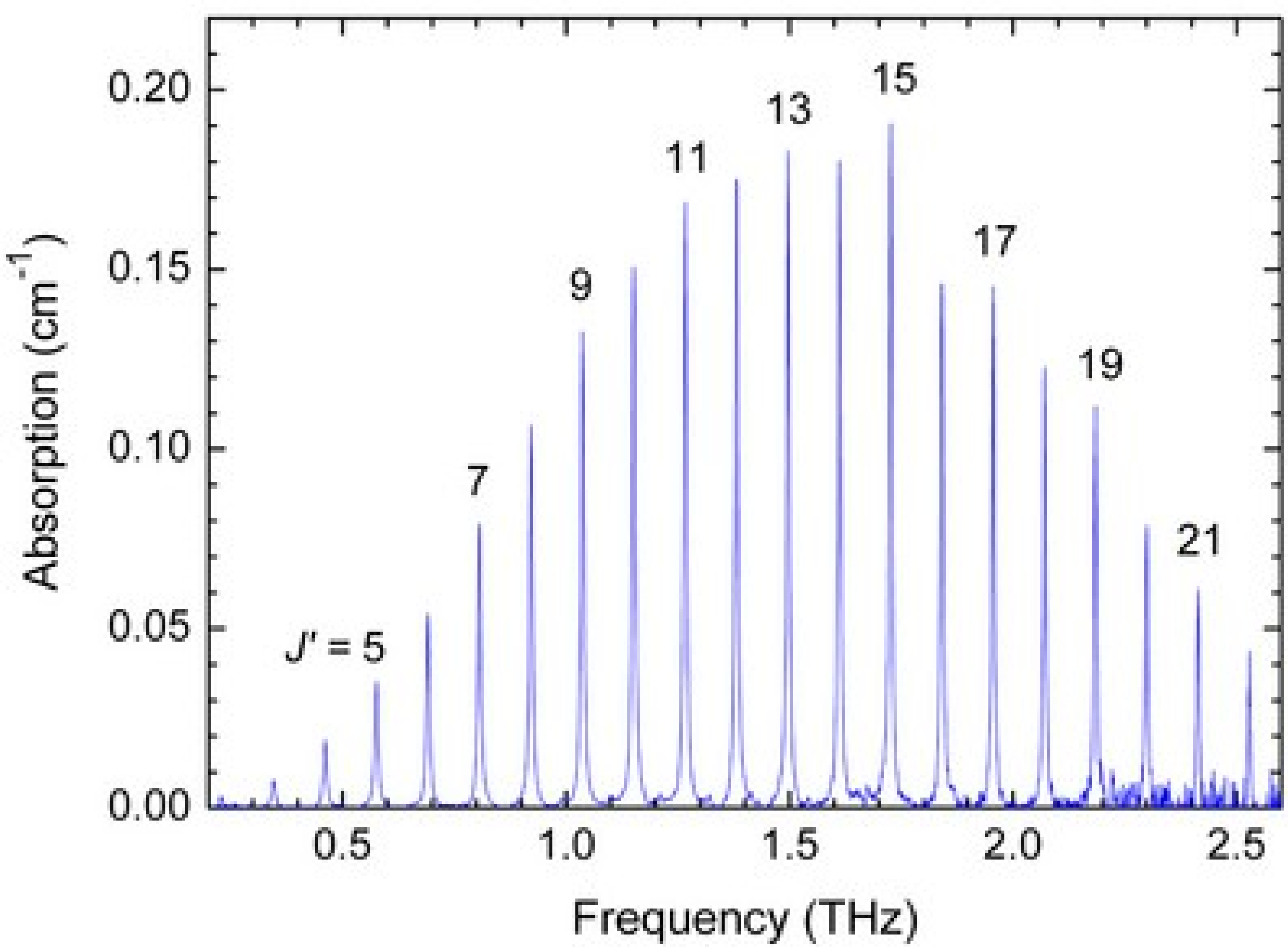
Brightness temperature related to the temperature of τ_{CO} surface

Firstly detect at least three lines (more is better!)

Line ratios determine T_{ex}

(which should be the same if Emission is optically thin)

Terahertz transmission spectrum of a CO cell at 2 bar pressure



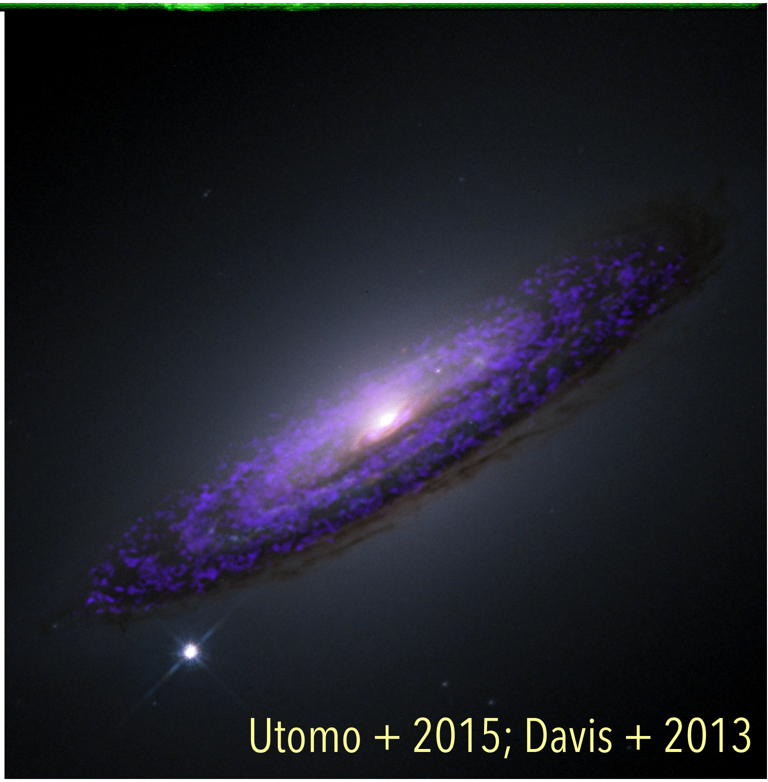
X_{CO} factor here better written

How to find the H_2 abundance: the X_{CO} factor:

$$X_{CO} = 1.6 \cdot 10^{19} \frac{T_{ex}}{30K} e^{\frac{5.5K}{T_{ex}} - 0.184} \quad cm^{-2} (K\,km\,s^{-1})^{-1}$$

A black hole mass measurement from molecular gas kinematics in NGC4526

We recently published a paper in the journal "Nature" where we measured the mass of a super-massive black hole (in galaxy NGC4526) by tracing the motions of molecular gas clouds swirling around it (see Figure 1, next page). This technique is exciting, because it opens up the possibility of measuring black hole masses in more galaxies than ever before. The research paper describing this work has been published in Nature (Febuary 15th 2013).



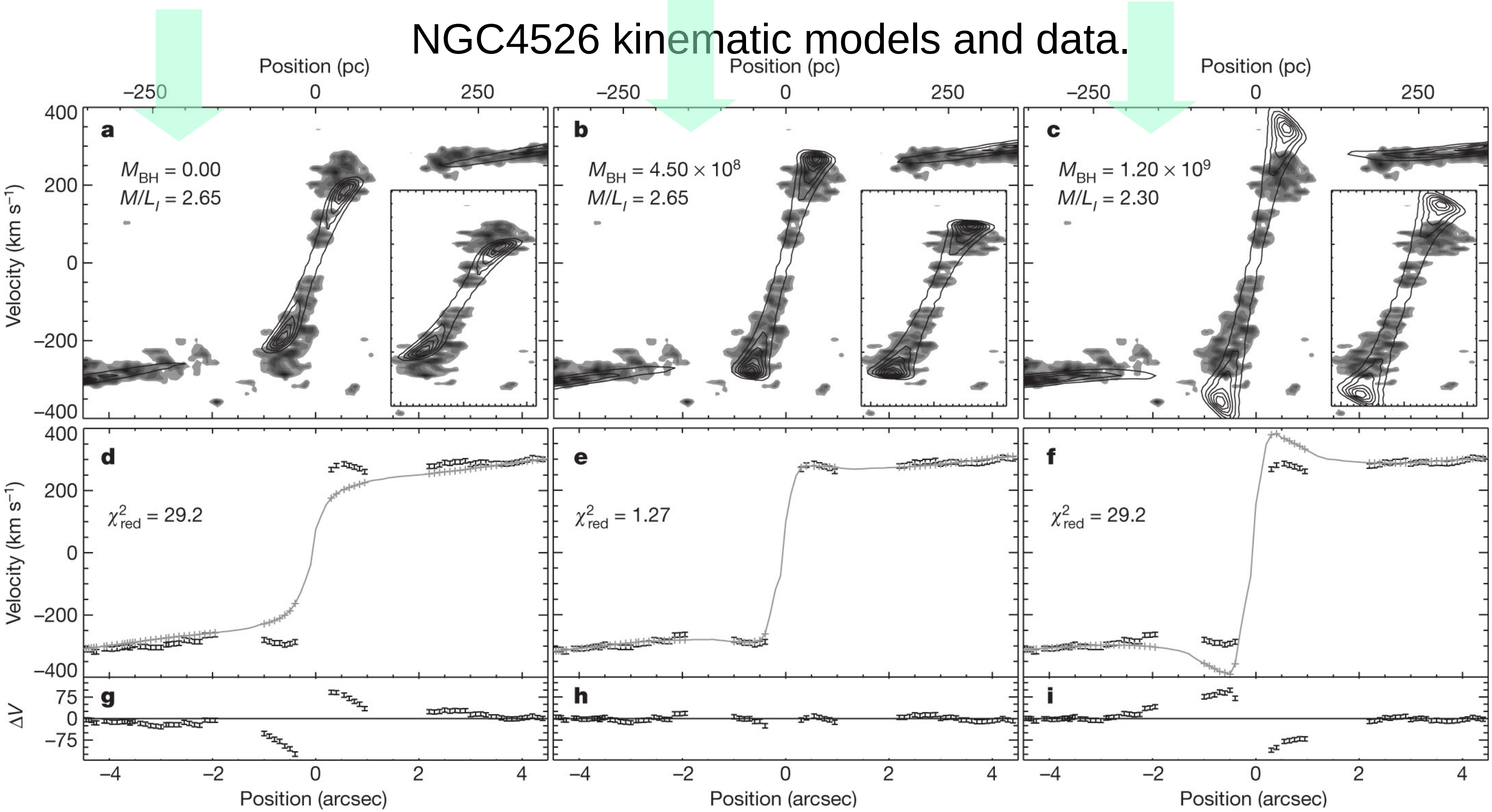
Super-Massive Black Holes

The masses of the super-massive black holes in the centres of galaxies appear to correlate with various properties of the parent galaxy (such as how bright the galaxy is, or the speed of random stellar motions). These relations extend from the biggest galaxies to small globular clusters, suggesting that there is a fundamental link between galaxy and black hole evolution. This is quite suprising, and not well understood, as these relations tie together quantities that probe very different length- and mass-scales.

These black hole correlations are based on a small number of black hole mass estimates, and the vast majority of the measurements have been made with just three methods. To understand the physical origin of the the black hole scaling relations we need more black hole mass measurements, and alternative techniques for measuring them.

The current methods for estimating black hole masses only work for relatively nearby galaxies, as you need to zoom in close to the black hole to measure its mass. Our paper demonstrates a new method for measuring black hole masses, with which we will be able to measure black hole masses much further out in the universe.

NGC4526 kinematic models and data.



TA Davis *et al.* *Nature* **000**, 1-3 (2013) doi:10.1038/nature11819

nature

Molecular Gas Measurements

Molecules that we are familar with in every day life also exist in space (however often in quite different proportions to those found on earth). The poisonous gas carbon monoxide is the second most abundant molecule in the universe. If we concentrated the gases in space to be as dense as our atmosphere, the result would be very toxic indeed! !

In this Nature paper, we observed carbon monoxide molecules in the galaxy NGC4526. We used the Combined Array for Research in Millimetre-wave Astronomy (CARMA) telescope (which you can see in Figure 3). CARMA uses an array of small dishes to obtain much sharper images than a single telescope could alone. This technique is called interferometry, and you can learn more about it here.

We observed NGC4526 with CARMAs sharpest array, achieving a resolution of 0.25 arcseconds. This is the equivalent of being able to spot a one euro coin (or US quarter) being held up 10 kilometres away! With these super sharp images we were able to zoom right into the centre of NGC4526, and observe the gas wizzing around the black hole. See the map of carbon monoxide in NGC4526 in Figure 1.

Gas (or any object) that orbits a black hole moves with a the speed that is determined by the mass of the black hole, and the distance from it. Our observations with CARMA give us the velocity of the gas at each position. Using this we were able to `weigh' the black hole in the centre of NGC4526, determining its mass.

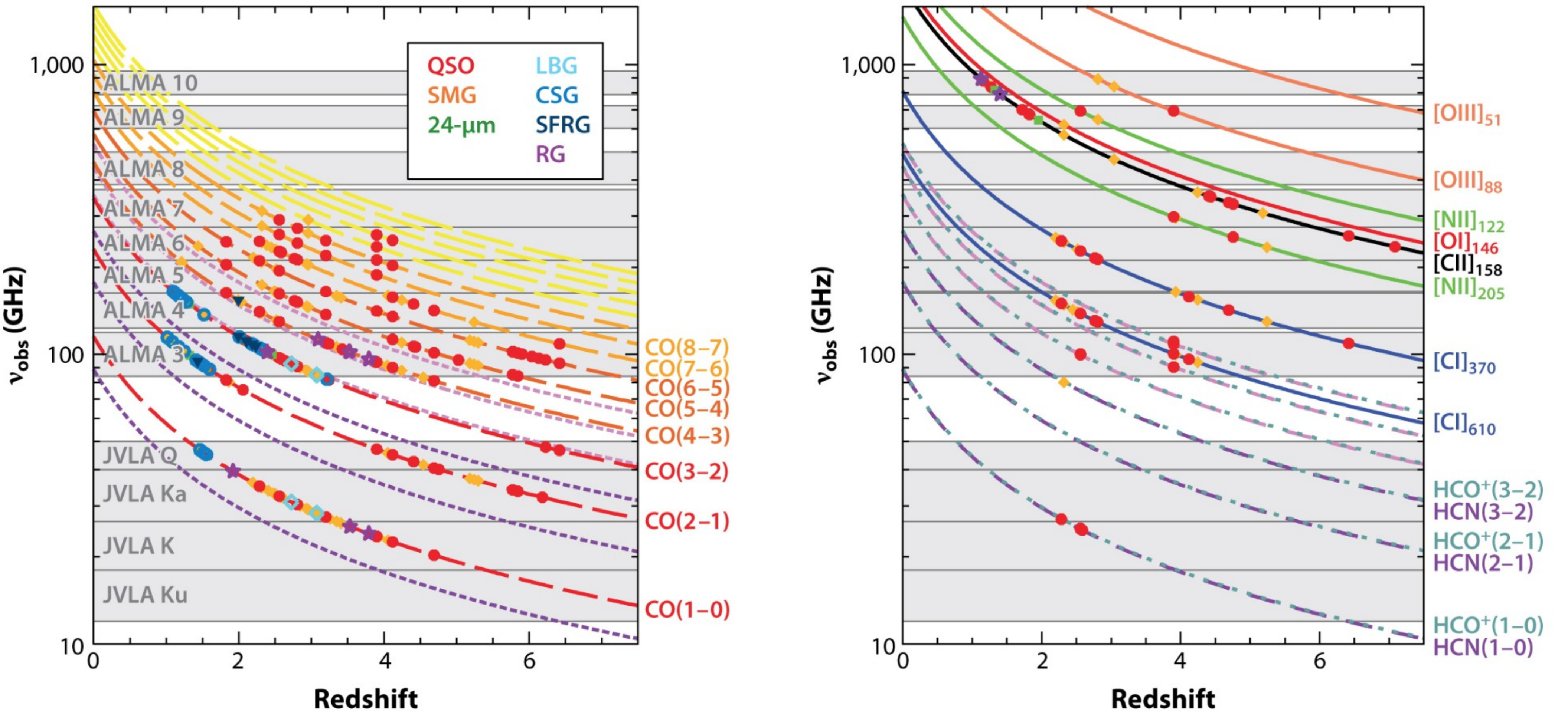
Incomplete list of known molecules

http://www.eana-net.eu/index.php?page=Wiki/interstellar_medium

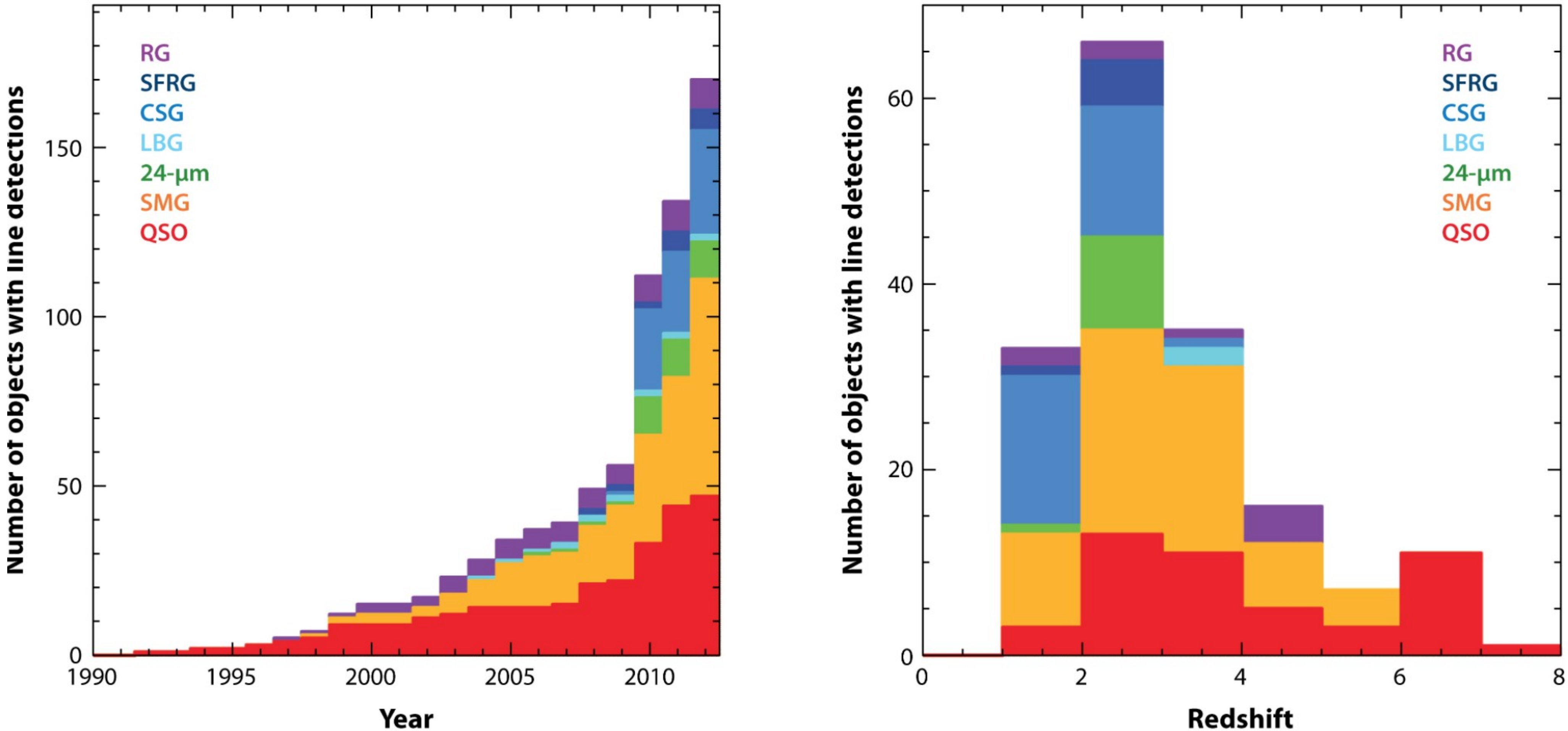
https://science.gsfc.nasa.gov/691/cosmicice/interstellar.html

Number of Atoms										
2	3	4	5	6	7	8	9	10	11	12+
H_2	C_3	c-C_3H	C_5	C_5H	C_6H	CH_3C_3N	CH_3C_4H	CH_3C_5N?	HC_9N	C_6H_6
AlF	C_2H	l-C_3H	C_4H	l-H_2C_4	CH_2CHCN	$HCOOCH_3$	CH_3CH_2CN	$(CH_3)_2CO$		$HC_{11}N$
AlCl	C_2O	C_3N	C_4Si	C_2H_4	CH_3C_2H	CH_3COOH?	$(CH_3)_2O$	NH_2CH_2COOH?		PAHs
C_2	C_2S	C_3O	l-C_3H_2	CH_3CN	HC_5N	C_7H	CH_3CH_2OH			C_{60}^{+}?
CH	CH_2	C_3S	c-C_3H_2	CH_3NC	$HCOCH_3$	H_2C_6	HC_7N			
CH^+	HCN	C_2H_2	CH_2CN	CH_3OH	NH_2CH_3	$HOCH_2CHO$	C_8H			
CN	HCO	CH_2D^+?	CH_4	CH_3SH	c-C_2H_4O					
CO	HCO^+	HCCN	HC_3N	HC_3NH^+						
CO^+	HCS^+	$HCNH^+$	HC_2NC	HC_2CHO						
CP	HOC^+	HNCO	HCOOH	NH_2CHO						
CSi	H_2O	HNCS	H_2CHN	C_5N						
HCl	H_2S	$HOCO^+$	H_2C_2O							
KCl	HNC	H_2CO	H_2NCN							
NH	HNO	H_2CN	HNC_3							
NO	MgCN	H_2CS	SiH_4							
NS	MgNC	H_3O^+	H_2COH^+							
NaCl	N_2H^+	NH_3								
OH	N_2O	SiC_3								
PN	NaCN	CH_3								
SO	OCS									
SO^+	SO_2									
SiN	c-SiC_2									
SiO	CO_2									
SiS	NH_2									
CS	H_3^+									
HF	H_2D^+									

Tracing reservoirs of cold gas in



Carilli CL, Walter F. 2013.
Annu. Rev. Astron. Astrophys. 51:105–61



Carilli CL, Walter F. 2013.
Annu. Rev. Astron. Astrophys. 51:105–61

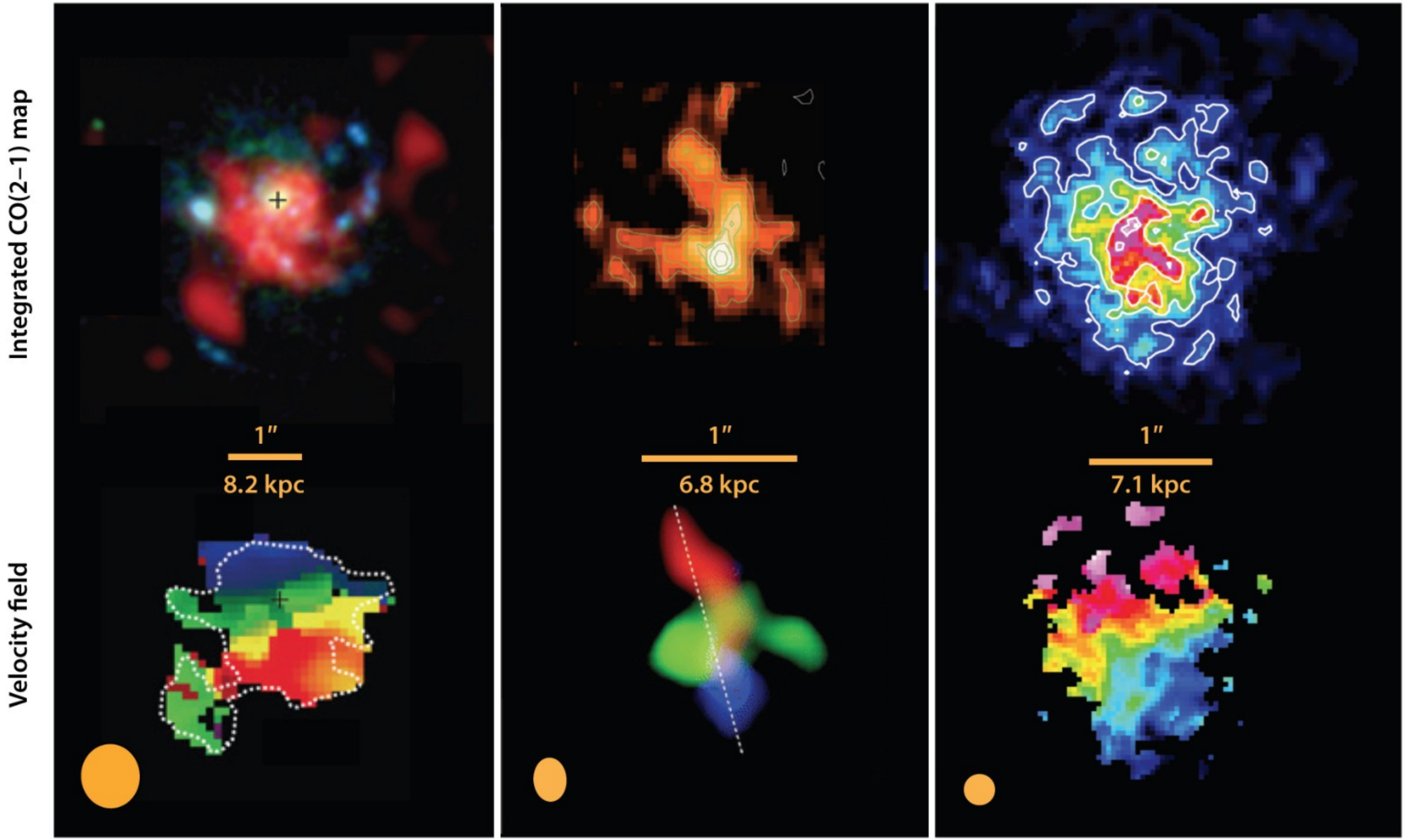
Tracing reservoirs of cold gas in (proto –) galaxies

Fuel for star formation and AGN activity (DIFFERENT SCALES!)
Star formation: --> infrared & UV emission, + radio, X-rays, etc
AGN: --> various phenomena: (optical) emission lines, UV & X–Rays, IR, radio emission, γ– rays

Observed in several hundreds Galaxies at z > 1 (out to Z ~7 in AGN)

Gas is common in highly star-forming "sub–mm galaxies", particularly 1.5 < z < 2.5
Is there a variation of X_{CO} .vs. z and .vs. galaxy type ?

Observational requirements:
Lines move in frequency (z dependent!), instruments become redshift windows



Carilli CL, Walter F. 2013.
Annu. Rev. Astron. Astrophys. 51:105–61

Cold (molecular) gas is found in form of molecules associated with DUST (helps formation, shields from high energy radiation)

- *Homologous molecules do not have electric dipole, high energy transitions in very particular physical conditions (high T, high P), i.e. rare!*
- *CO, despite its low electric dipole is commonly seen, being the second to H_2 in abundance At typical ISM conditions J1 $\rightarrow$ 0 is weaker that higher (J $\rightarrow$ J-1) transitions, and it is optically thick when n> 2000 cm^{-3}*
- *Gas in all phases is found in Spiral & Irregular galaxies, most evident in interacting objects Thin disk distribution, avoids the bulge, and the outer parts (density is too low!)*
- *Little or no cold gas is found in elliptical and S0 galaxies, with a few exceptions*
- *Gas condensations create conditions for (intense) star formation*

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