Dynamic simulations of the Galactic Dynamo
based on supernova-driven ISM turbulence

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MKSP meeting, Sant'Antioco

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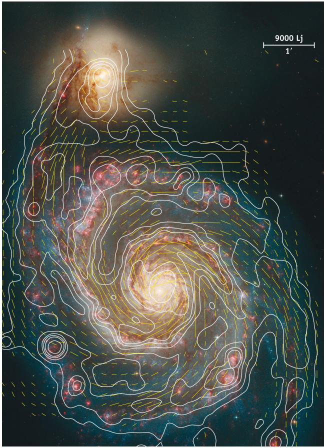
observations



Andrew Fletcher/Rainer Beck,
SuW and Hubble Heritage Team, STScI/AURA

- What is the origin of regular galactic magnetic fields?
  - primordial field, (i.e. frozen-in fossil record of galaxy formation)
  - dynamo-generated field, (i.e. dynamically replenished)
- Beck of the envelope
  - galactic rotation winds-up B_ϕ
    $\tau_\Omega \simeq 2\pi/25\,\mathrm{kpc}^{-1}\,\mathrm{km\,s}^{-1} \simeq 250\,\mathrm{Myr}$
  - turbulent diffusion
    $\tau_\mathrm{d} \simeq (0.5\,\mathrm{kpc})^2/0.5\,\mathrm{kpc\,km\,s}^{-1} \simeq 500\,\mathrm{Myr}$
  - large observed pitch angle strongly favours dynamo

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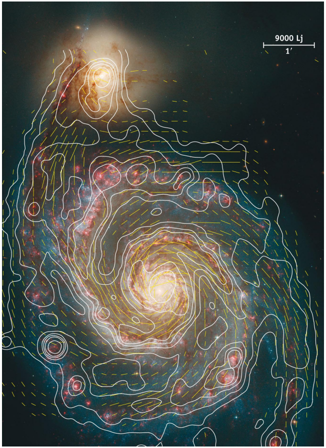

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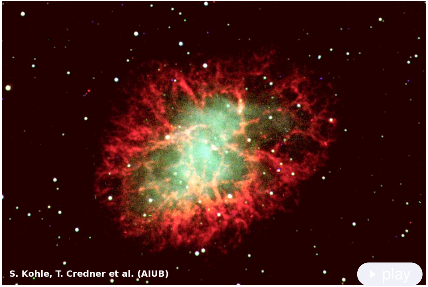
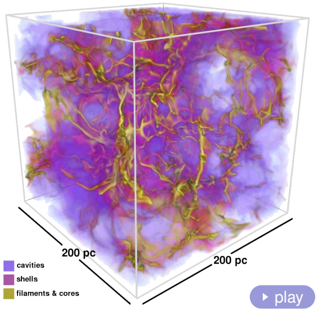
supernova-driven turbulence



- interstellar medium highly turbulent
- energy deposited by supernovae, CRs, MRI, stellar winds, protostellar jets, ...
- 2-3 SNe per century in our own Milky Way

- small-scale dynamo is simple (in a way...)
- but how amplify regular fields in a turbulent environment? rotation + stratification → mean-field dynamo

supernova-driven turbulence



play

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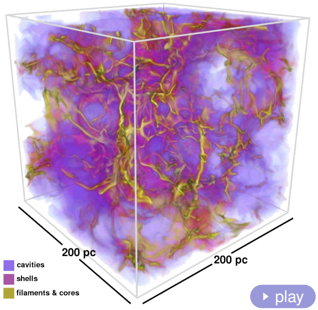

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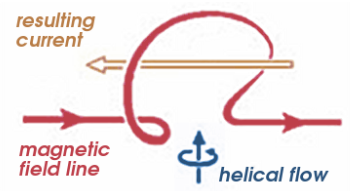
the α effect dynamo

- helical α effect in a single expanding SNR ▸ play
- breaking the homogeneity of the turbulence ▸ play

- Key mechanisms
  - rotation (and/or shear) $\rightarrow$ field-line stretching
  - helical flow component $\rightarrow$ avoid cancellation due to anti-parallel field
  - α effect couples the poloidal and toroidal field components
  - reconnection $\rightarrow$ restore original field-line topology

the α effect dynamo



- helical α effect in a single expanding SNR play
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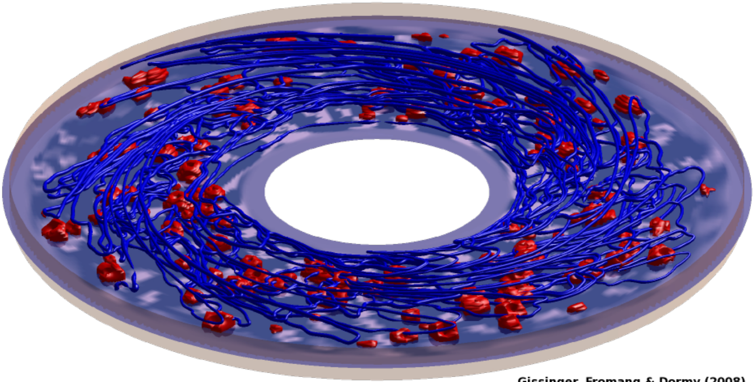
the big picture



Gissinger, Fromang & Dormy (2008)

- encapsulate the effect of the supernovae
- model the evolution of the large-scale field

modelling the dynamo process

- Mean-field approach:
  - split into mean + fluctuation

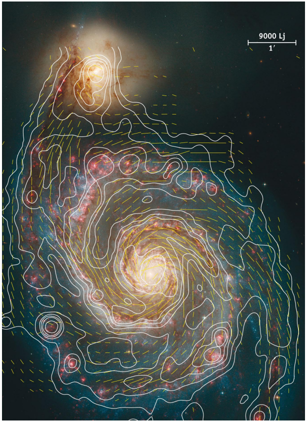
    $\mathbf{U} = \overline{\mathbf{U}} + \mathbf{u}$ and $\mathbf{B} = \overline{\mathbf{B}} + \mathbf{b}$
  - derive mean-field equation

$$\partial_t \overline{\mathbf{B}} = \nabla\times\left(\overline{\mathbf{U}}\times\overline{\mathbf{B}} + \bar{\mathcal{E}} - \eta\,\nabla\times\overline{\mathbf{B}}\right)$$

turbulent EMF $\bar{\mathcal{E}} = \overline{\mathbf{u}\times\mathbf{b}}$

- Parametrise small-scale effects $\bar{\mathcal{E}}$ as a functional of $\overline{\mathbf{U}}$, $\overline{\mathbf{B}}$, $\overline{f(\mathbf{u})}$
  - for sufficient scale separation

$$\bar{\mathcal{E}}_i = \alpha_{ij}\bar{B}_j - \tilde{\eta}_{ij}\,\varepsilon_{jkl}\partial_k\bar{B}_l$$



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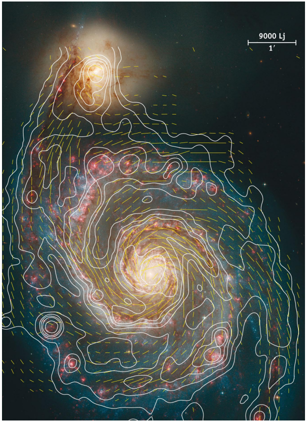

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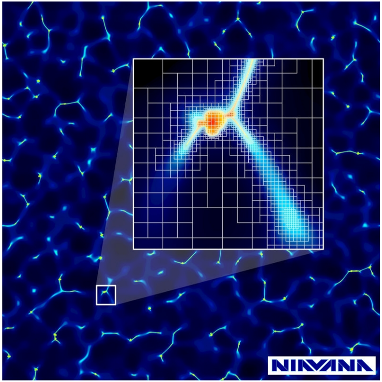
the nirvana code



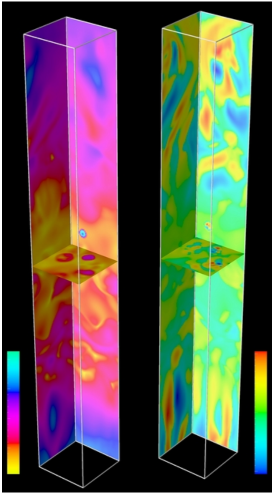
- Udo Ziegler's NIRVANA-III
  - conservative (dual energy) MHD grid code
  - 2nd-order central scheme + constrained transport (CT)
  - 3rd-order Runge-Kutta time integrator
  - block structured adaptive mesh refinement (AMR)
  - efficiently MPI-parallel (space-filling-curve techniques for AMR)

local box simulations

- Model geometry:
  - local patch of interstellar medium, up to 1.6 kpc on edge ($\Delta \simeq 10$ pc)
  - vertical stratification up to ± 6 kpc
  - sheared galactic rotation
- Physical ingredients:
  - non-ideal MHD (+ heat conduction)
  - optically thin radiative heating/cooling
  - localised thermal energy input modelling the supernovae

play

Korpi, Brandenburg, Shukurov, Tuominen & Nordlund (1999)



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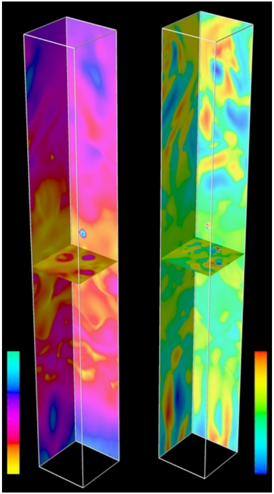
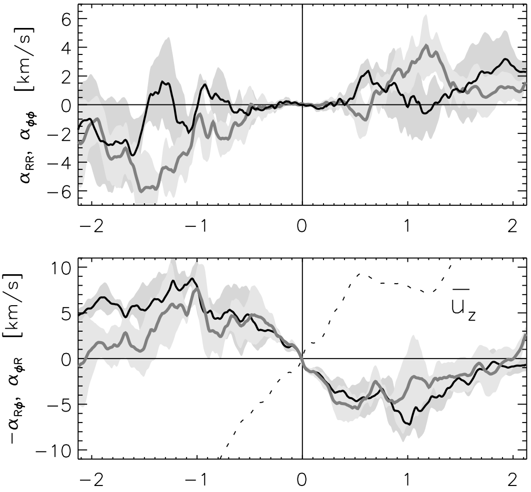


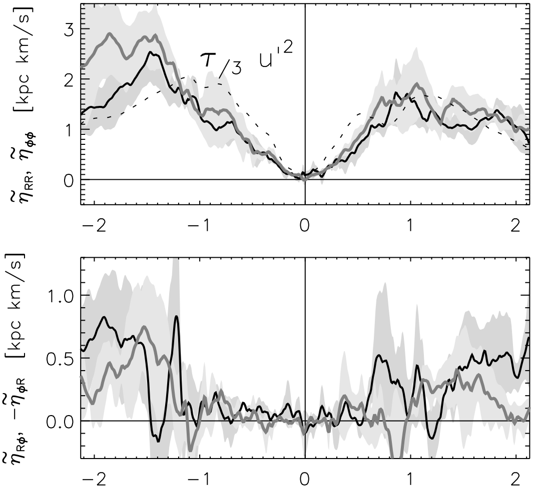
α-profiles



- dynamo effect $|\alpha_R|, |\alpha_\phi| \simeq 3\,\mathrm{km\,s^{-1}}$
- diamagn. pumping $|\gamma_z| \simeq 7\,\mathrm{km\,s^{-1}}$ directed inward
- $|\alpha| : |\gamma|$ consistent w/ SOCA results
- effect of galactic wind $\bar{u}_z$ balanced by turb. pumping

Gressel, Elstner, Ziegler & Rüdiger (2008), A&A 486, L35

$\tilde{\eta}$-profiles



- turb. diffusivity $\simeq 2\,\mathrm{kpc\,km\,s^{-1}}$
- coherence time $\tau \simeq 3\,\mathrm{Myr}$

- non-vanishing $\Omega \times J$ effect $\delta_z \simeq 0.5\,\mathrm{kpc\,km\,s^{-1}}$
- add shear $\rightarrow$ dynamo

it really works!

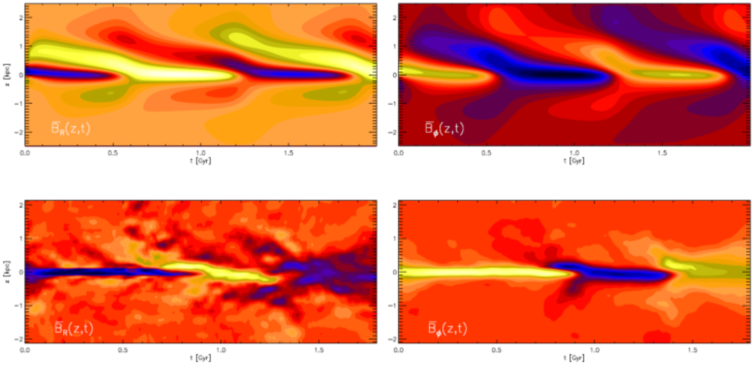


Figure 4.10: Same as Fig. 4.9, but additionally including a mixed (anti-)symmetric contribution in the off-diagonal elements of $\tilde{\eta}$ (upper panels). Now the lopsided dipolar symmetry in the field reversals persists and closely resembles the features seen in the *direct* simulation H4 (lower panels).

scaling relations (brute force)

- dynamo effect as function of
  - supernova rate $\hat{\sigma} = \sigma/\sigma_0$
  - rotation frequency $\hat{\Omega} = \Omega/\Omega_0$
  - midplane density $\hat{\rho} = \rho/\rho_0$

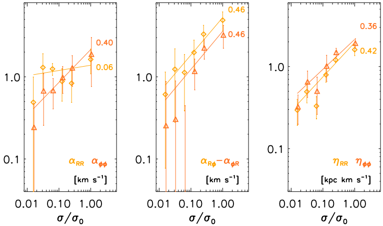
- scaling relations:

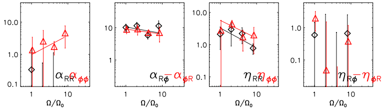
$$\alpha = 2\ \mathrm{km\,s^{-1}}\quad \hat{\sigma}^{0.4}\ \hat{\Omega}^{0.5}\ \hat{\rho}^{-0.1}$$

$$\gamma_z = 12\ \mathrm{km\,s^{-1}}\quad \hat{\sigma}^{0.5}\ \hat{\Omega}^{-0.2}\ \hat{\rho}^{0.3}$$

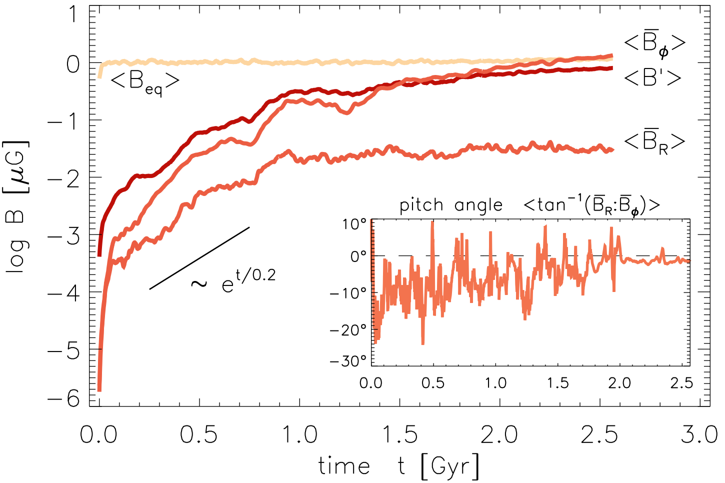
$$\eta = 2\ \mathrm{kpc\,km\,s^{-1}}\ \hat{\sigma}^{0.4}\ \hat{\Omega}^{0.25}\ \hat{\rho}^{0.4}$$

where $\sigma_0 = 30\mathrm{Myr}^{-1}\mathrm{kpc}^{-2}$,
$\Omega_0 = 25\mathrm{Gyr}^{-1}$, $\rho_0 = 1\,\mathrm{cm}^{-3}$

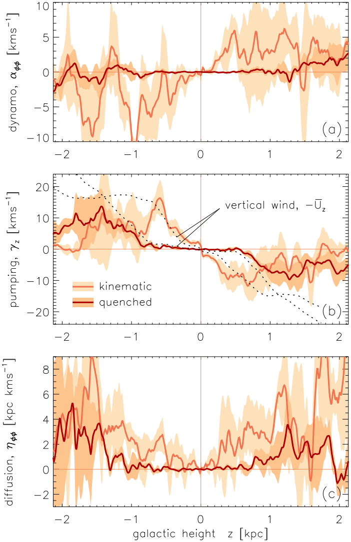




magnetic field saturation



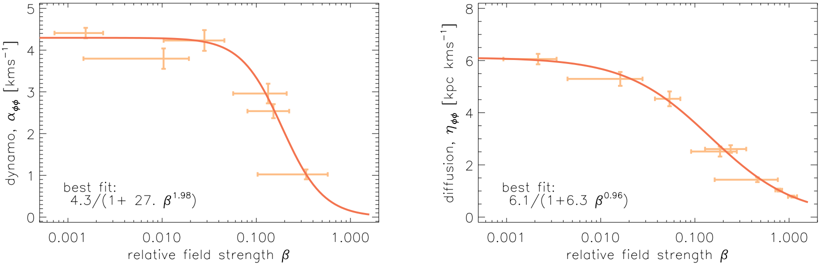
a lingering catastrophe



- Quenching scenarios:
  - (a) classic: flow quenching due to Lorentz force
  - (b) catastrophic: helicity conservation inhibits growth
  - (c) similar to scenario (b) but alleviated by small-scale helicity removal
- Test possible realisations:
  - quenching sets-in . . .
    (a) . . . at $B \simeq B_{\rm eq}$
    (b) . . . at $B \simeq B_{\rm eq}/{\rm Rm}$
    (c) . . . at $B \simeq B_{\rm eq}\, l_0/L_0$
- Suppression of wind: (c) $\rightarrow$ (b)

extracting quenching functions

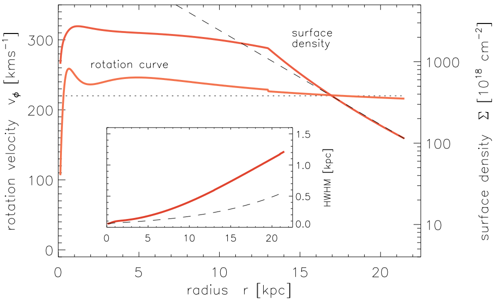
- quenching quadratic in $\beta \equiv \bar{B}/B_{\rm eq}$
- magnetic Reynolds number, $\mathrm{Rm} \equiv u_{\rm rms}(k_{\rm f}\,\eta)^{-1} \simeq 75$–$125$
- scale separation ratio, $l_0/L_0 \simeq 0.1\,\mathrm{kpc}/1\,\mathrm{kpc} = 10$



Gressel, Bendre & Elstner (2013), MNRAS 429, 967

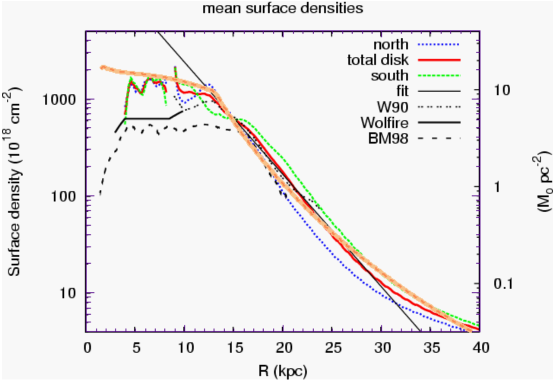
the disc model

- Based on flaring HI disc
  Kalberla & Dedes (2008)
  - expon. + power-law density profile
  - NFW-type DM halo + stellar disc / bulge
  - $\rightarrow$ self-consistent rotation curve



- Goal: perform fully-dynamical MHD + MFD simulations
  - momentum equation with turbulent viscosity
  - will capture Parker / Tayler / MRI on long wavelengths

disc surface density



Kalberla & Dedes (2008)

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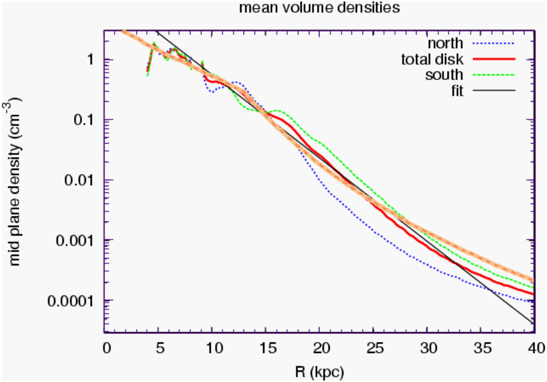
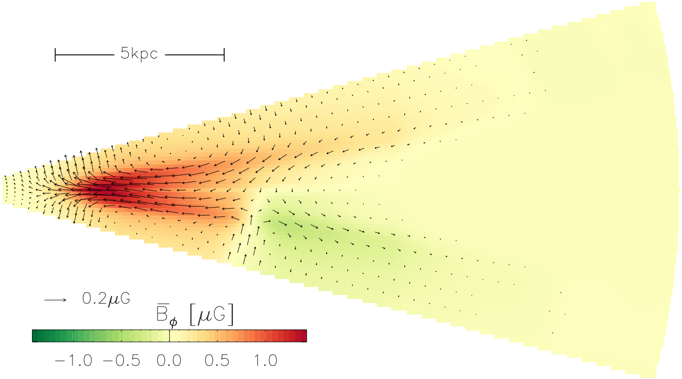


Kalberla & Dedes (2008)

emerging dynamo mode



- Dynamo solution
  - initial NVF leads to transient A0, → S0 mode
  - S0+A0 produces one-sided vertical field   Mao, Gaensler et al. (2010)

model overview

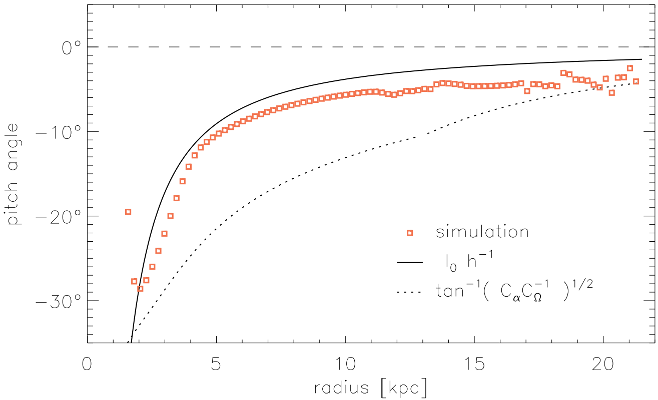
Table 2. Simulation parameters and results.

model	dim	MF	NS	halo	M_{gas} [10^{10}M$_\odot$]	seed	parity	p_{in} [°]	p_{out} [°]	τ_e [Gyr]	$\lvert\overline{B}_{sat}\rvert$ [μG]	comments
X1s-0.5	2D	•	○	○	0.57	WN	S0/A0	-31.2	-4.7	0.374	1.44	see Fig. 3
X1s	2D	•	○	○	1.14	WN	S0[a]	-28.7	-4.6	0.503	3.75	see Figs. 4,5
X1s-1.5	2D	•	○	○	1.70	WN	S0[a]	-32.7	-4.6	0.547	6.34	
X1s-2.0	2D	•	○	○	2.27	WN	S0	-35.8	-4.3	0.593	9.07	
X2s-halo	2D	•	○	•	1.14	WN	S0[a]	-25.2	-4.8	0.358	4.05	
X3s-VF	2D	•	○	○	1.14	NVF	A0→S0	-28.9	-4.5	0.539	3.75	
N1s/d-HF	3D	•	○	○	1.14	NHF	S0	-28.7	-3.2	–	3.75	
	3D	•	•	○	1.14		S0	-17.7	-2.6	–	2.52	
N1s/d-VF	3D	•	○	○	1.14	NVF, B_ϕ	S0	-29.0	-6.1	0.409	3.75	
	3D	•	•	○	1.14		S0	-17.8	-2.7	0.407	2.65	see Fig. 8
N2d	3D	○[b]	•	○	1.14	HF+VF	A0	-5.3	-1.6	–	0.82[c]	see Fig. 6
N2d-MRI	3D	○	•	○	1.14	HF+VF	A0	-6.0	-0.5	–	4.25	see Fig. 7

[a] sub-dominant A0 outside $R \simeq 10$ kpc, [b] includes η_t, and ν_t, [c] obtained outside $R \simeq 15$ kpc.

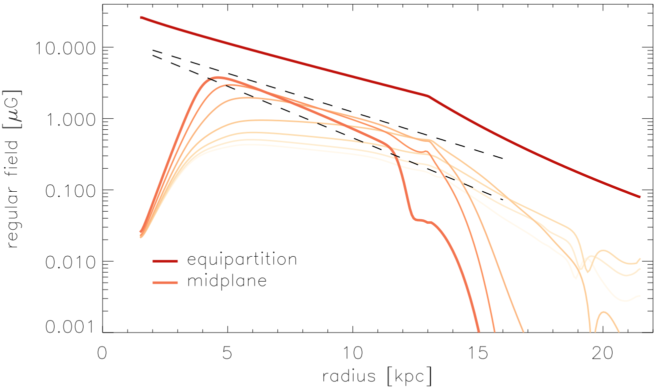
All 2D runs are axisymmetric; mean-field (MF) effects include the ones described in Sect. 2.2; runs including 'NS' evolve the Navier-Stokes equation. The 'halo' dynamo is shown as a dashed line in Fig. 2. The column labelled M_{gas} gives the normalisation for the disc mass. For seed fields we use white noise (WN), net-vertical field (NVF), net-horizontal field (NHF). Pitch angles are given for the inner disc (peak value) and for the outer disc (average for $R > 10$ kpc) separately. Growth rates are for the magnetic field $\lvert\overline{B}\rvert$, during an interval for which exponential growth can be identified.

radial pitch angle



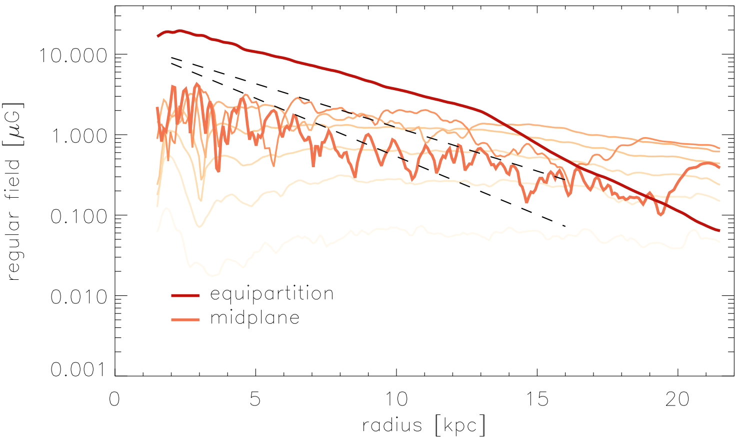
- radial fall-off in pitch angle (agrees with observations)
- → explained conveniently by flaring disc   Fletcher (2010)

saturated field profile



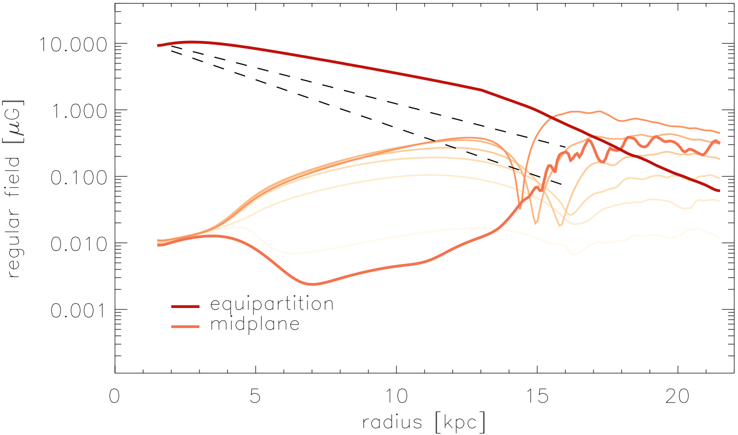
- radial scale length $\sim 4\,\mathrm{kpc}$ for saturated $\overline{\mathbf{B}}$
- outer disc essentially unmagnetised

MHD w/o mean-field dynamo



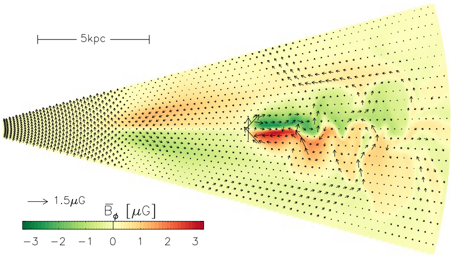
MRI-only simulation MRI with dissipation from SNe

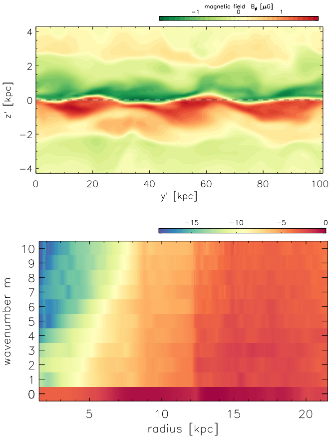
MHD w/o mean-field dynamo



MRI-only simulation MRI with dissipation from SNe

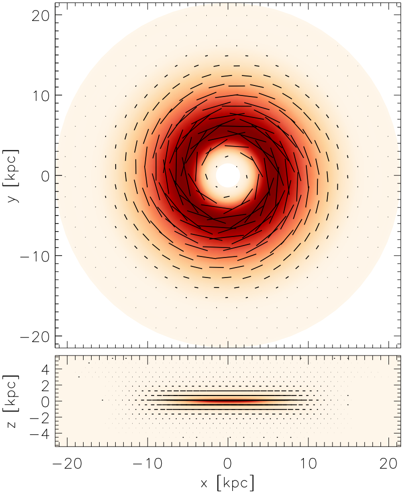
combined MHD / mean-field dynamo

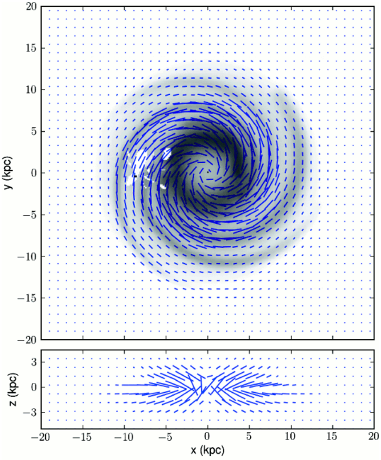




- Parker modes & MRI in outer disc
  $\rightarrow$ pronounced loop structures
- undulating mode with $m = 3$
- radial scale length $\sim 10\,\text{kpc}$
- inner disc dominated by $m = 0$

synthetic polarisation maps





Jansson & Farrar (2012)

summary of results

- Measuring dynamo coefficients
  - 1D mean-field model matches simulations
  - $\rightarrow$ quantitative scaling relations for sub-grid physics
- Non-linear saturation
  - quenching functions obtained
  - indications for the presence of helicity constraints
  - suppression of wind threatens saturation level
- Global mean-field models
  - first fully quantitative global dynamo models
  - dynamic momentum equation $\rightarrow$ MRI / Parker / Tayler
  - parametrisation of small-scale effects is essential

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Developing Tools for Synthetic Polarization Maps
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