What Sets the Maximum Power of Jets from Black Holes?

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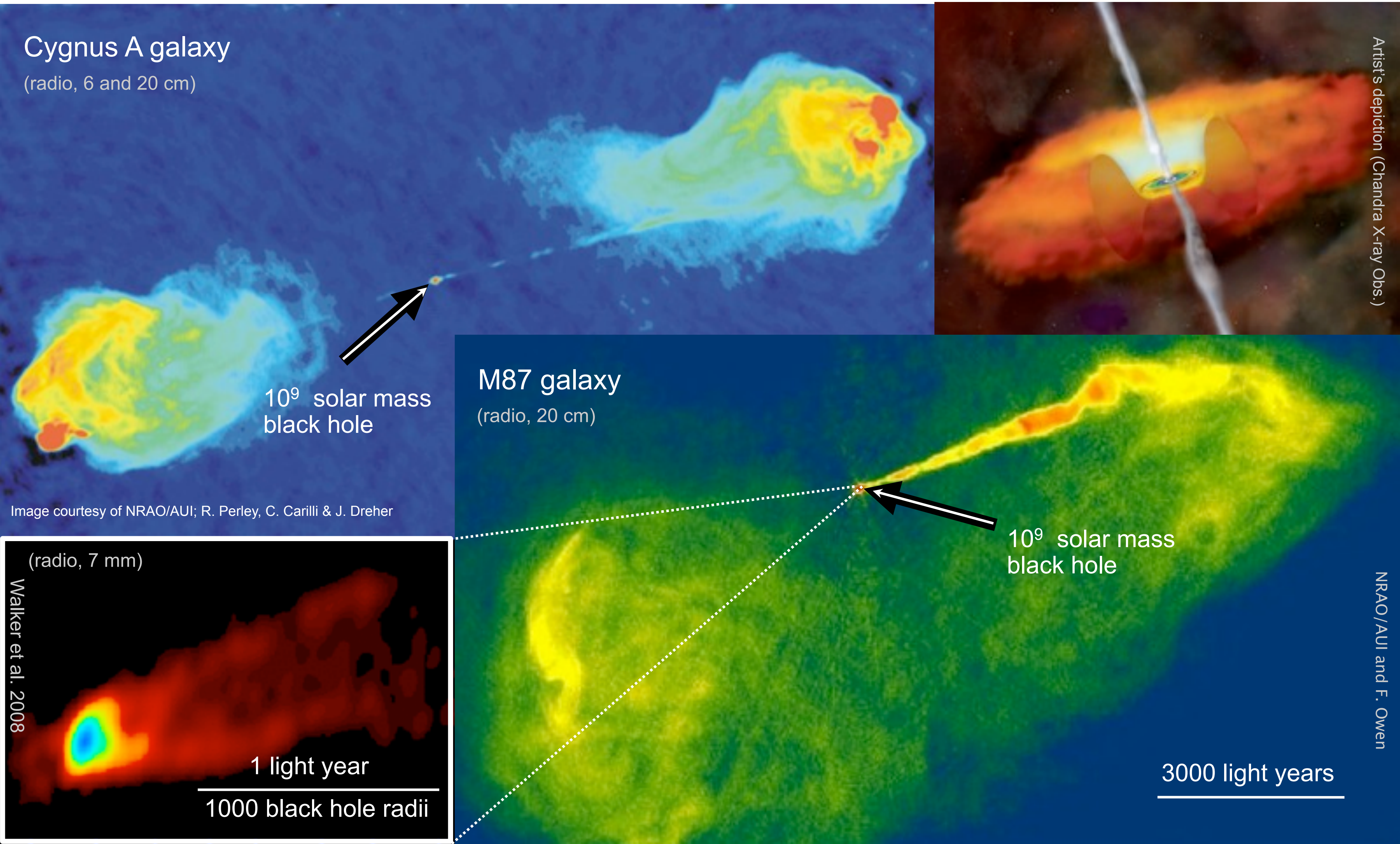
Jets: Beautiful & Challenging



Fundamental Questions

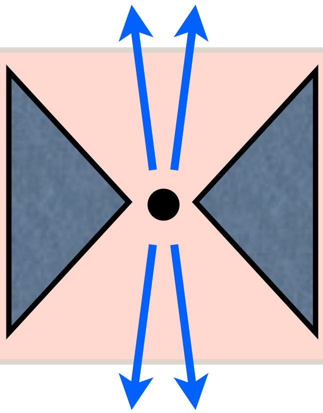
- What sets the maximum power of jets?
- Are jets powered by black holes or inner regions of accretion disks?
- How does jet power depend on BH spin?
- Does accretion always spin up BHs to high spins?

When Do We Observe Jets?

$\lambda = L/L_{\rm edd}$

$h/r \sim 1$

$\tau \gg 1$



Radiatively-Inefficient
(super-Eddington)

100%

10%

$h/r \ll 1$

$\tau \gg 1$

Thin Disk

(High/Soft or Thermal state)

1%

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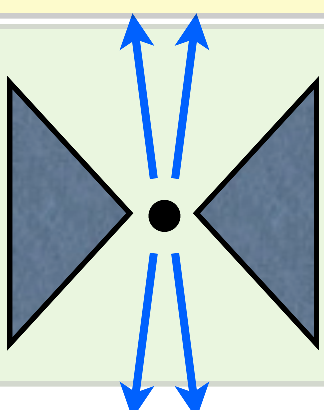

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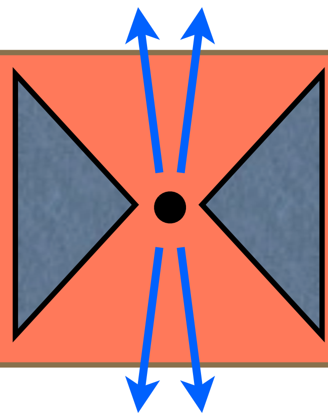

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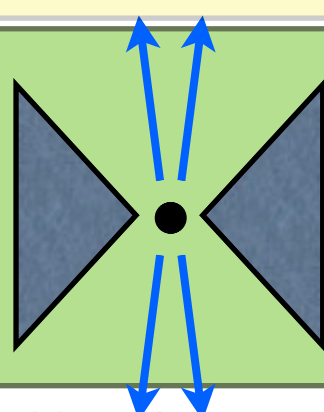

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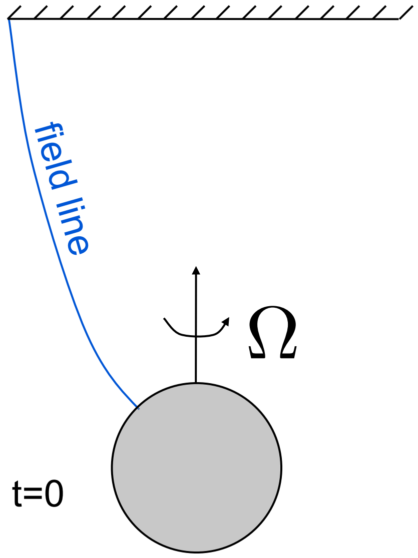
How Do Jets Work?



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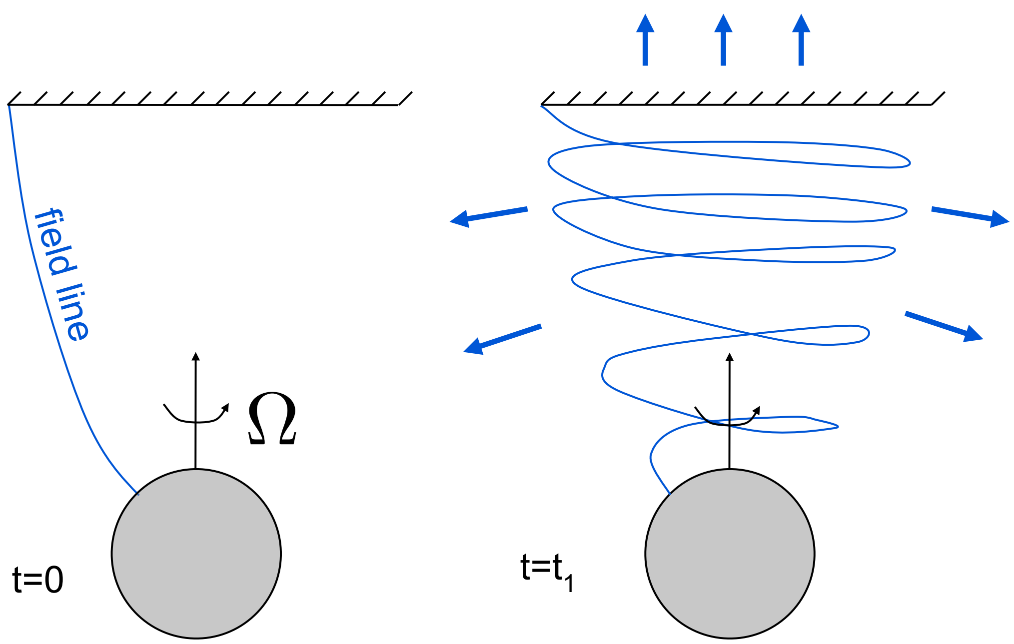
$$P = \frac{B_\varphi^2}{8\pi}$$



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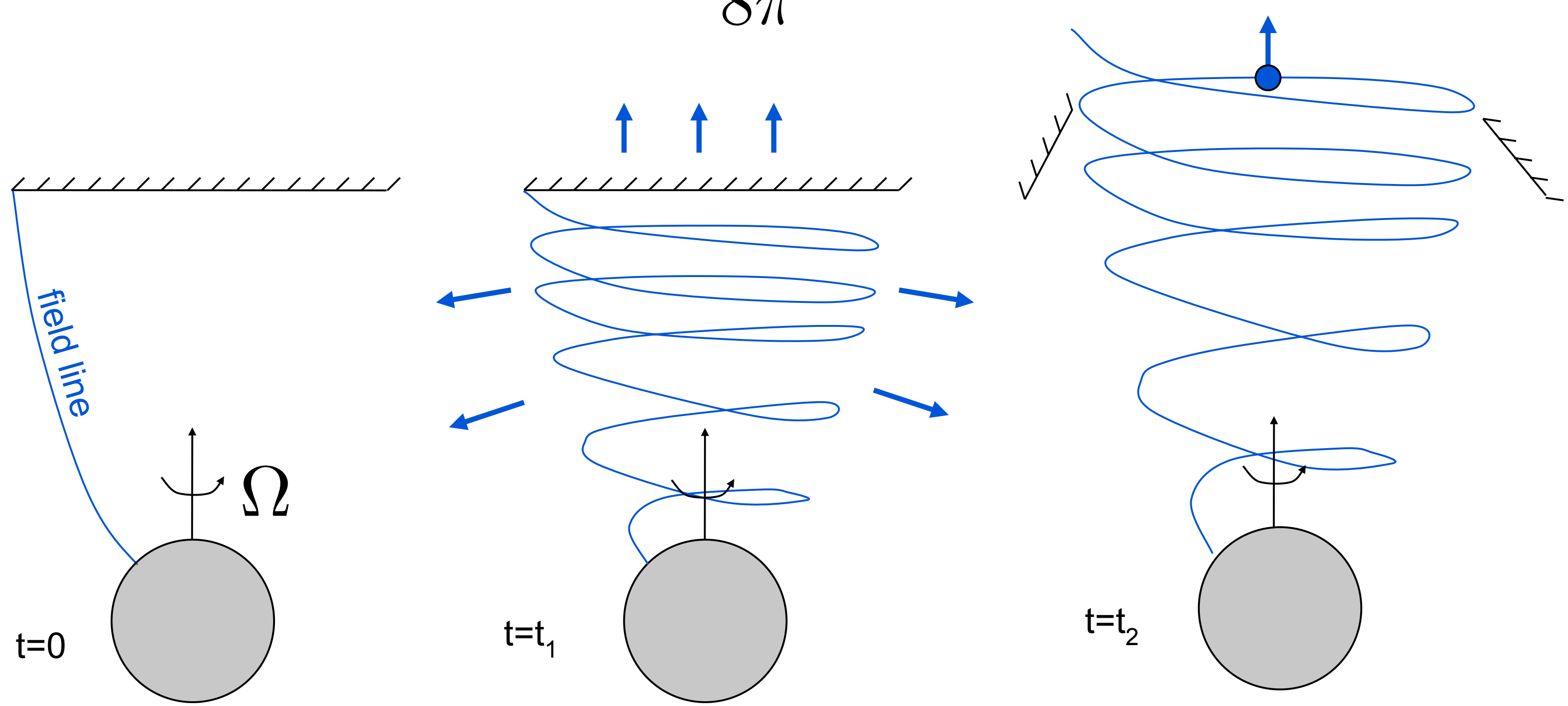
Field toroidally-dominated $B_\varphi \gg B_z$



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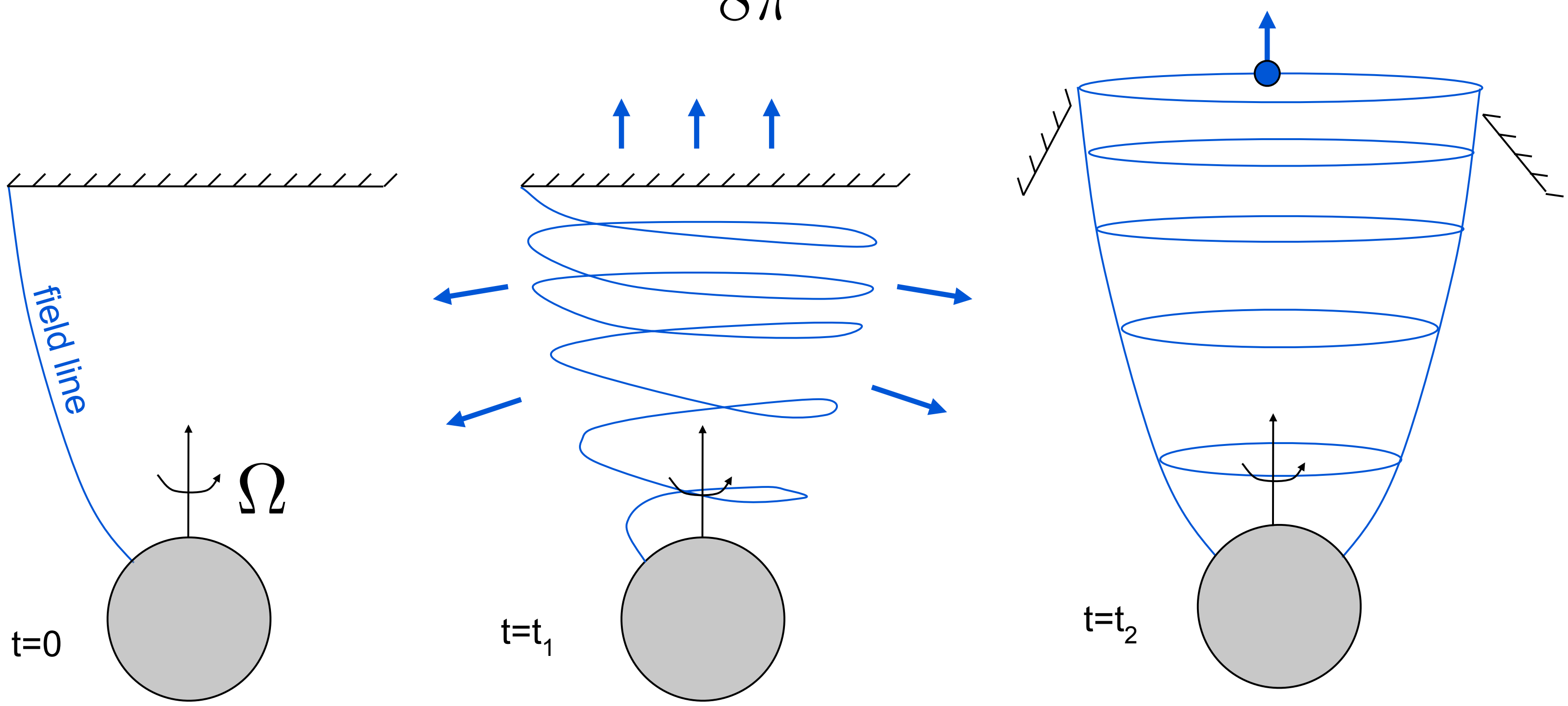

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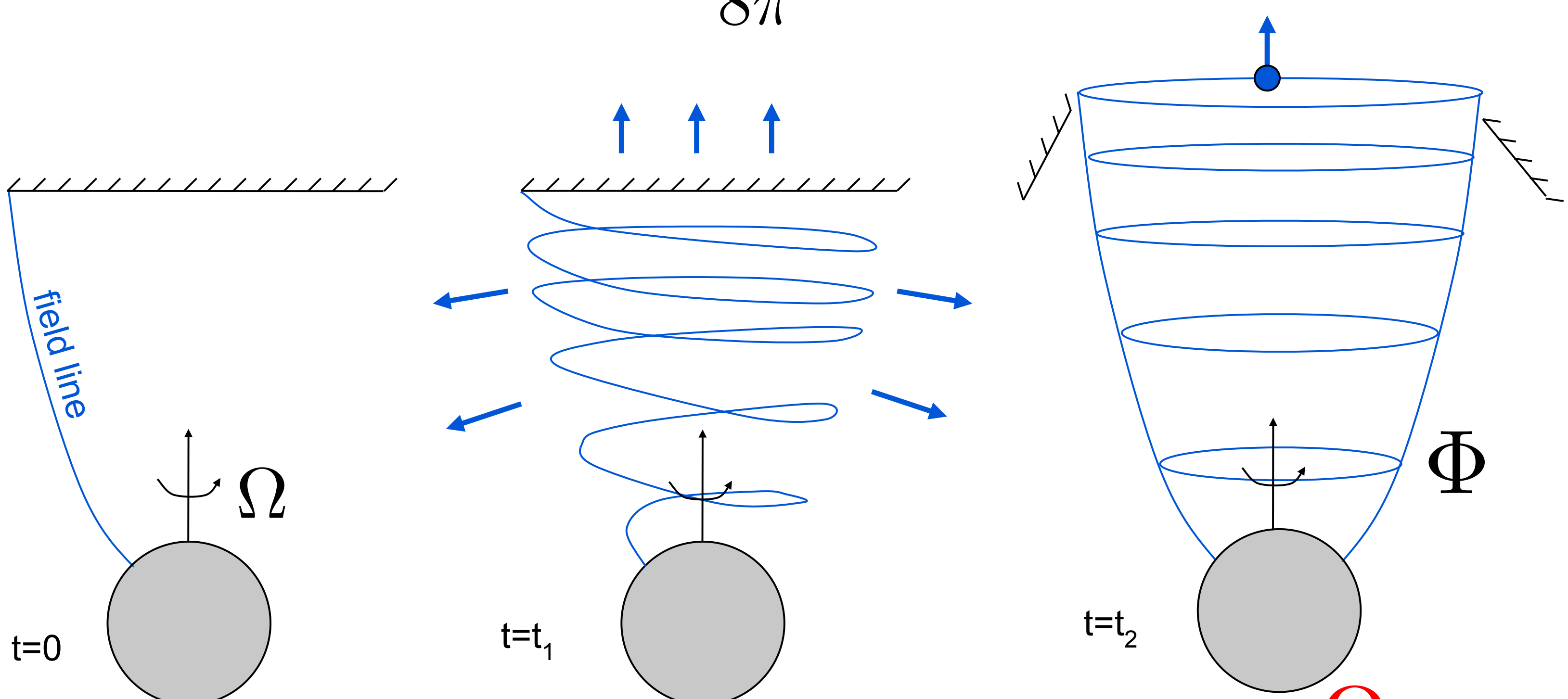

$$\Omega_{\rm H} = \frac{ac}{2r_{\rm H}}$$

What Sets BH Power?

- We understand well how BH power depends on Φ and $\Omega_{\rm H}$:

$$P_{\rm BZ} = \frac{k}{c}\Phi^2\Omega_{\rm H}^2$$

(Blandford & Znajek 1977,
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- But, what sets value of the proportionality factor,

$$\phi = \frac{\Phi}{\sqrt{\dot{M} r_g^2 c}},$$

(Gammie 2002, Komissarov & Barkov 2009, Penna et al. 2010)

and BH power efficiency,

$$\eta_{\rm BZ} = \frac{P_{\rm BZ}}{\dot{M}c^2} = k\,\phi^2 \left(\frac{\Omega_{\rm H} r_g}{c}\right)^2 \times f(\Omega_{\rm H})$$

What sets the flux?

- Gravity limits BH field strength:

$$\frac{B^2}{8\pi} \lesssim \frac{GM\Sigma}{R^2}$$ (Narayan+ 2003)

- Mass conservation in a disk:

$$\dot{M} = 2\pi R \Sigma \beta_r c$$

- BH magnetic field:

$$B_r^{\mathrm{max}} \sim 2 \times 10^4 \ [\mathrm{G}]\, (0.1/\beta_r)^{1/2}$$

$$\phi^{\mathrm{max}} \simeq 50\, (0.1/\beta_r)^{1/2}$$

for $\dot{M} = 0.1 \dot{M}_{\mathrm{Edd}}, M = 10^9 M_{\odot}$

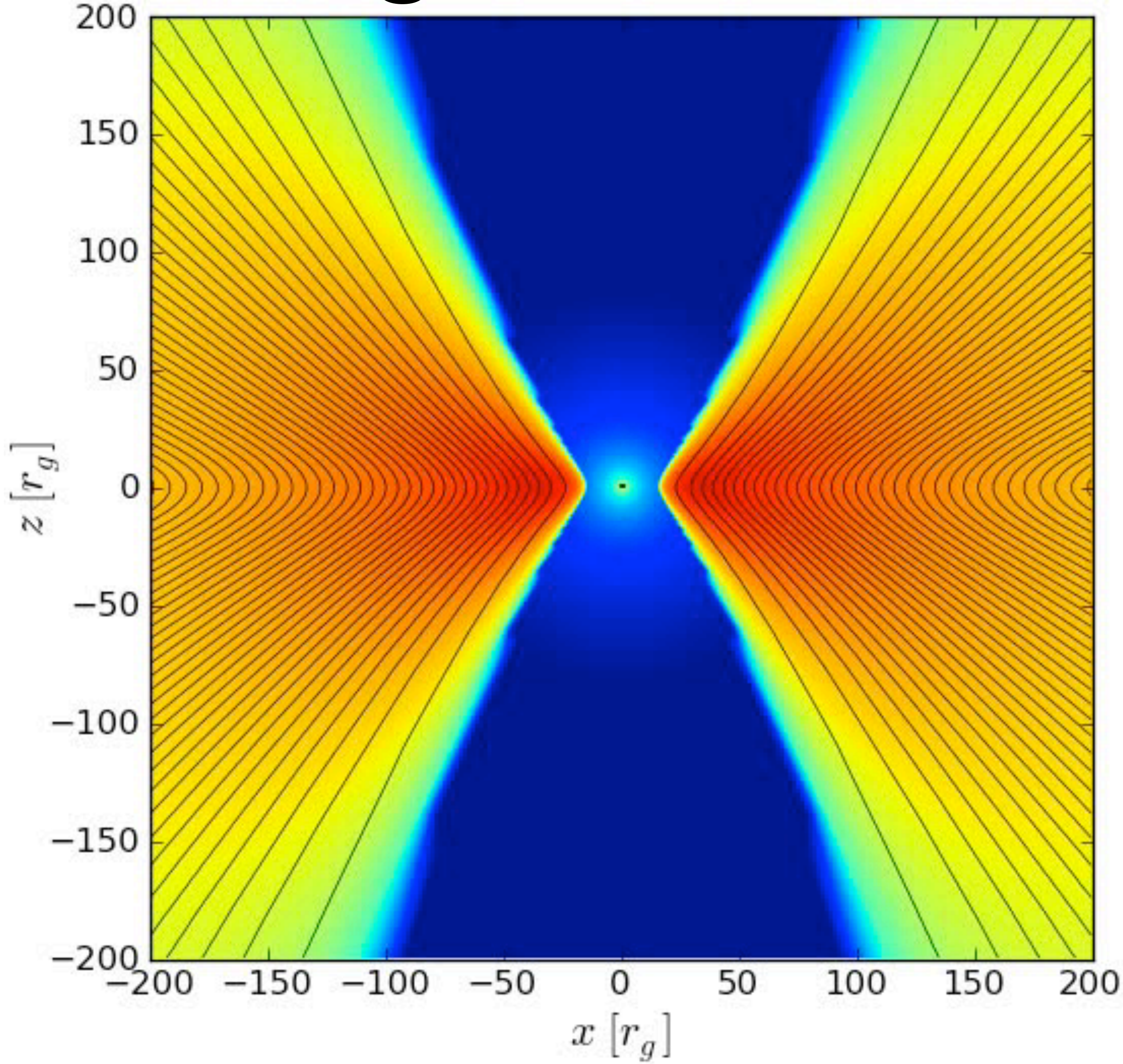
Maximum Black Hole Power = ?

- Jet power depends on magnetic field topology (McKinney 2005, Beckwith, Hawley & Krolik 2008, McKinney & Blandford 2009)
  - ➡ Dipolar geometry gives powerful jet
  - ➡ Quadrupolar or toroidal gives weak or no jet
- GR MHD simulations give $\eta_{\rm BZ} \lesssim 20\%$, even for nearly maximally spinning BHs (McKinney 2005, de Villiers et al. 2005, Hawley & Krolik 2006, Barkov & Baushev 2011)
- Can we obtain larger values of η?
- Observations: some AGN have $\eta \gtrsim 100\%$ (Rawlings & Saunders 1991, Fernandes et al. 2010, Ghisellini et al. 2010, Punsly 2011, McNamara et al. 2011)

Maximum Black Hole Power = MAD Power

- Jet power increases with increasing BH magnetic flux, Φ.
- BH + large Φ = magnetically-arrested accretion (MAD): BH is saturated with flux, and B-field is as strong as gravity (Bisnovatyi-Kogan & Ruzmaikin 74, 76, Igumenshchev et al. 03, Narayan et al. 03, AT et al. 11, AT & McKinney 12a,b, McKinney, AT, Blandford 12).
  - ➡ Maximum η for each spin
  - ➡ Efficiency exceeds 100%
  - ➡ First example of net energy extraction from a BH
- New physics: high jet power, QPOs, mode of accretion...
- Advanced 3D GR MHD simulations with HARM code: took over 2000 CPU-years! (Gammie et. al. 2003; AT et al. 2007)

Much Larger Flux than Before



Our grid extends out to $10^5 r_g$

AT, Narayan, McKinney 2011, MNRAS, 478, L79

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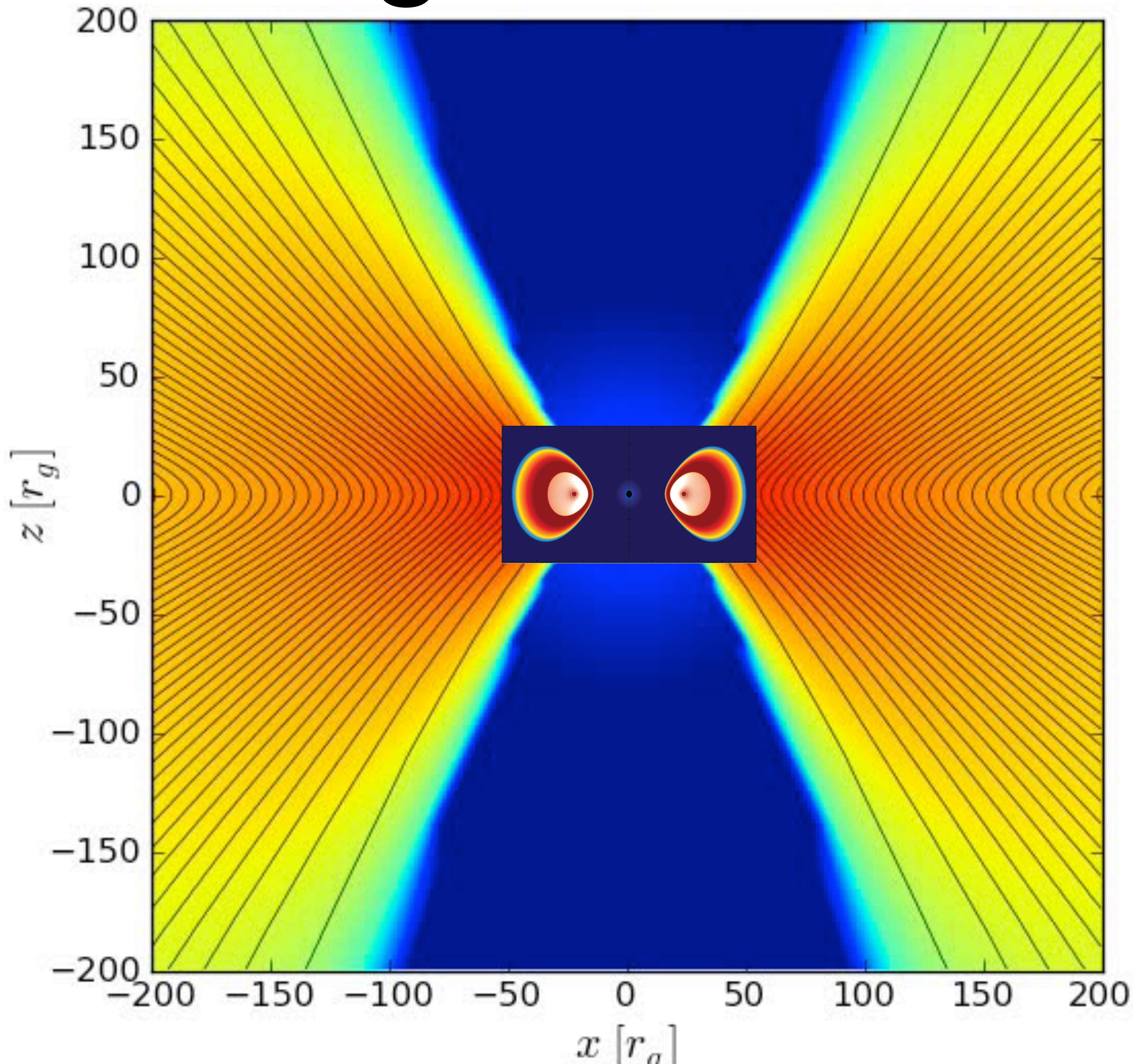

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Beckwith, Hawley, Krolik 2008

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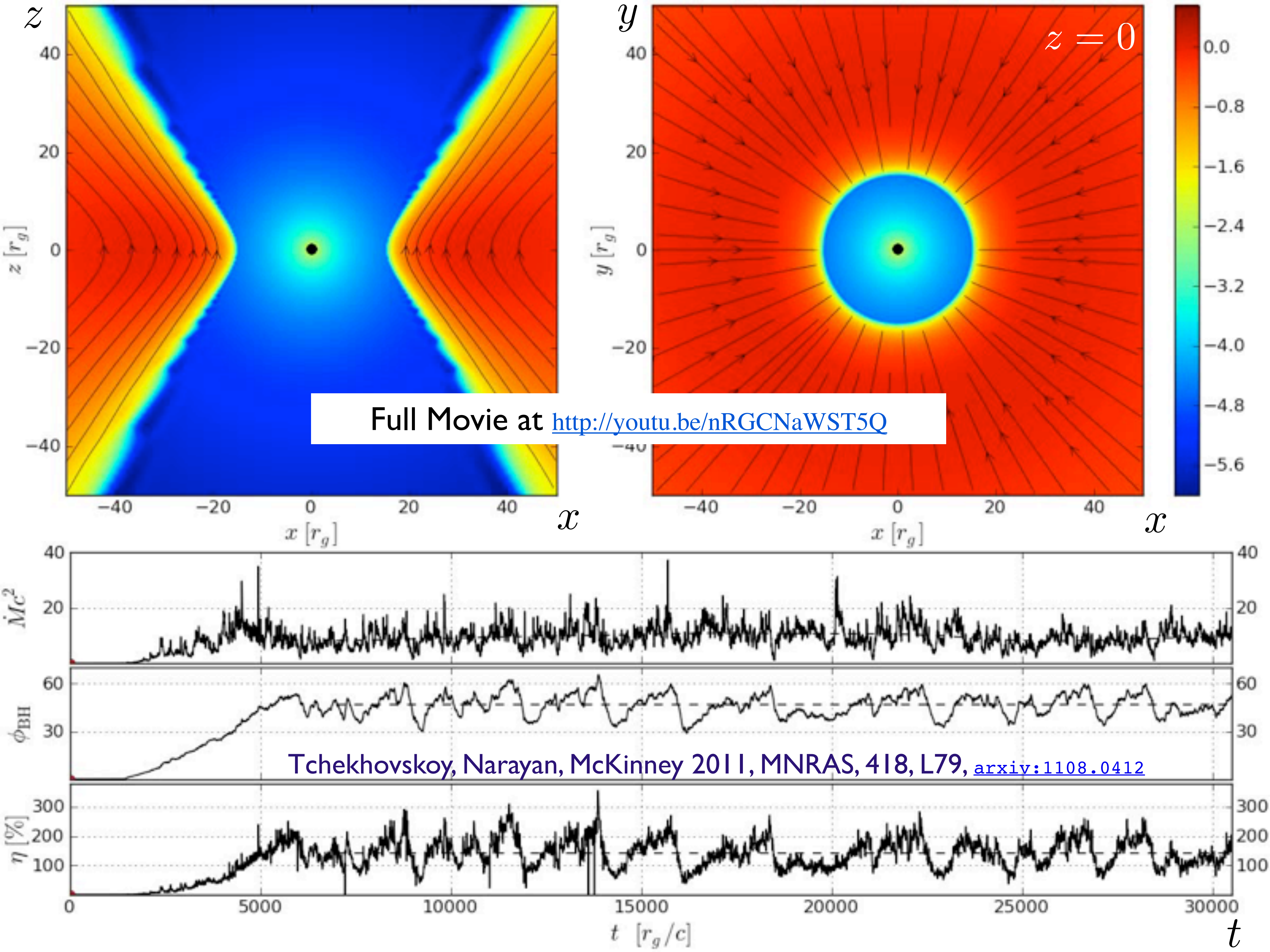
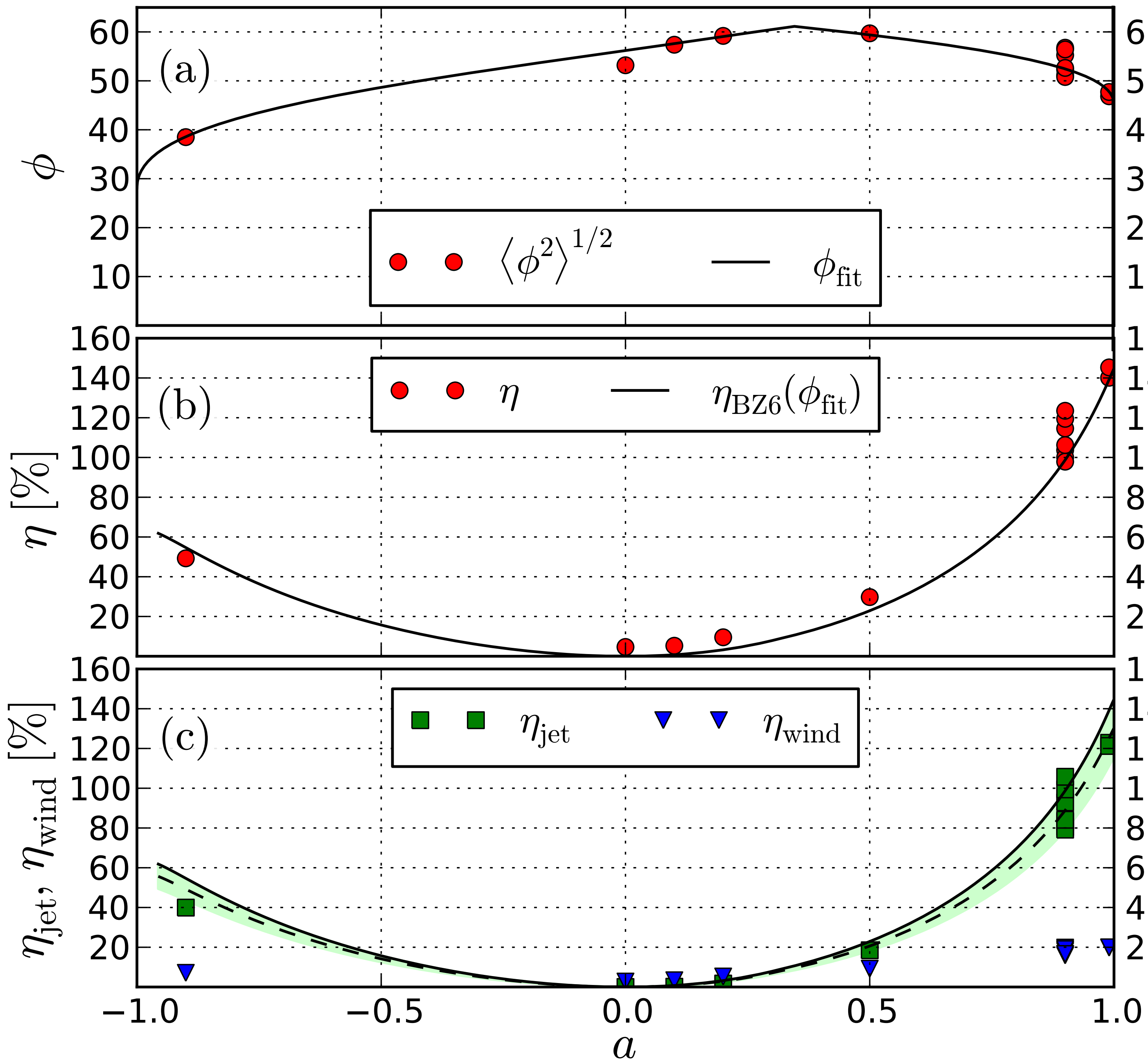


Full Movie at http://youtu.be/nRGCNaWST5Q

Tchekhovskoy, Narayan, McKinney 2011, MNRAS, 418, L79, arxiv:1108.0412

Maximum jet power vs. spin (h/r~0.3)



(AT, McKinney 2012a, MNRAS, accepted, arxiv:1201.4385 2012b, in prep.)

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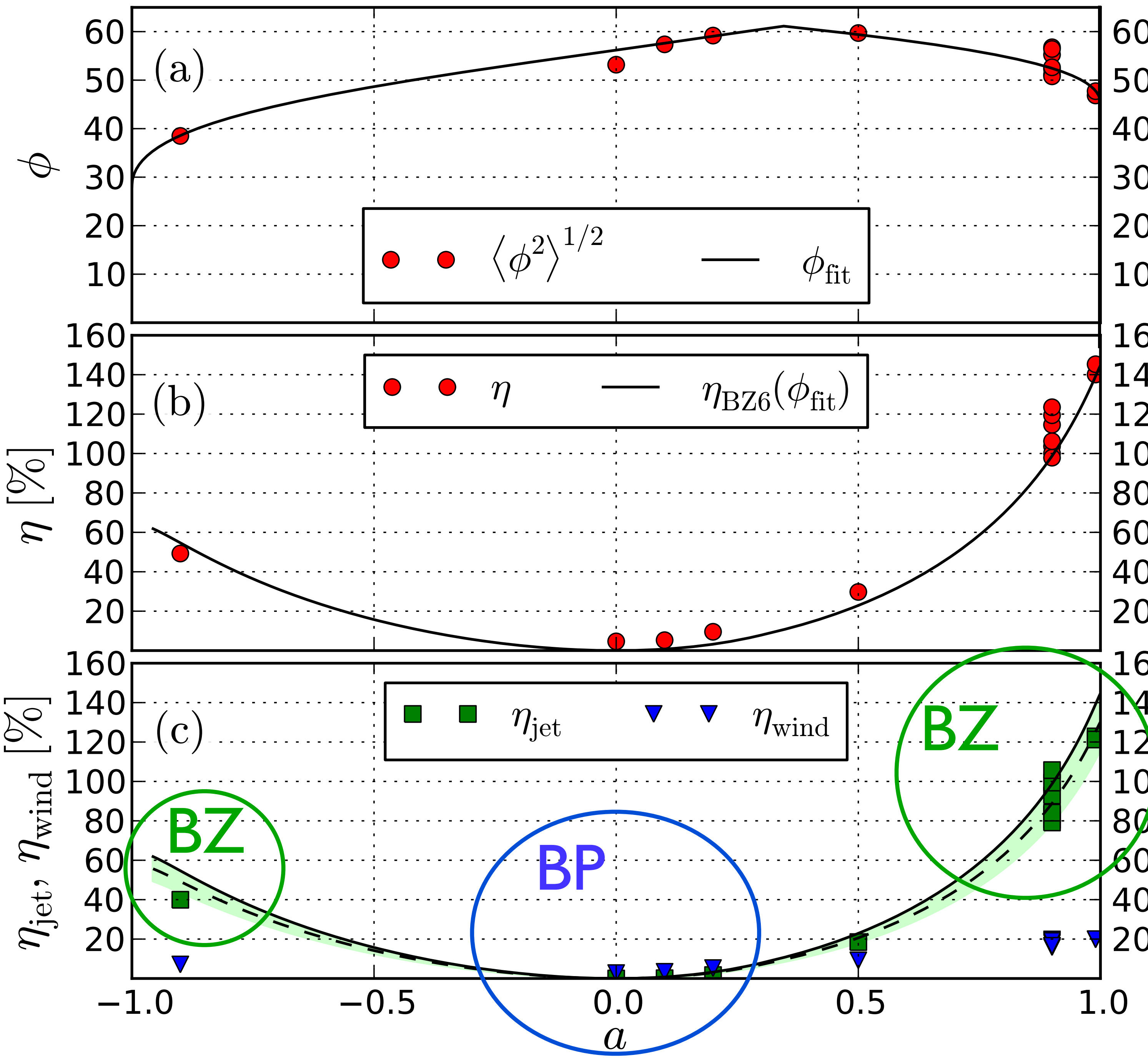

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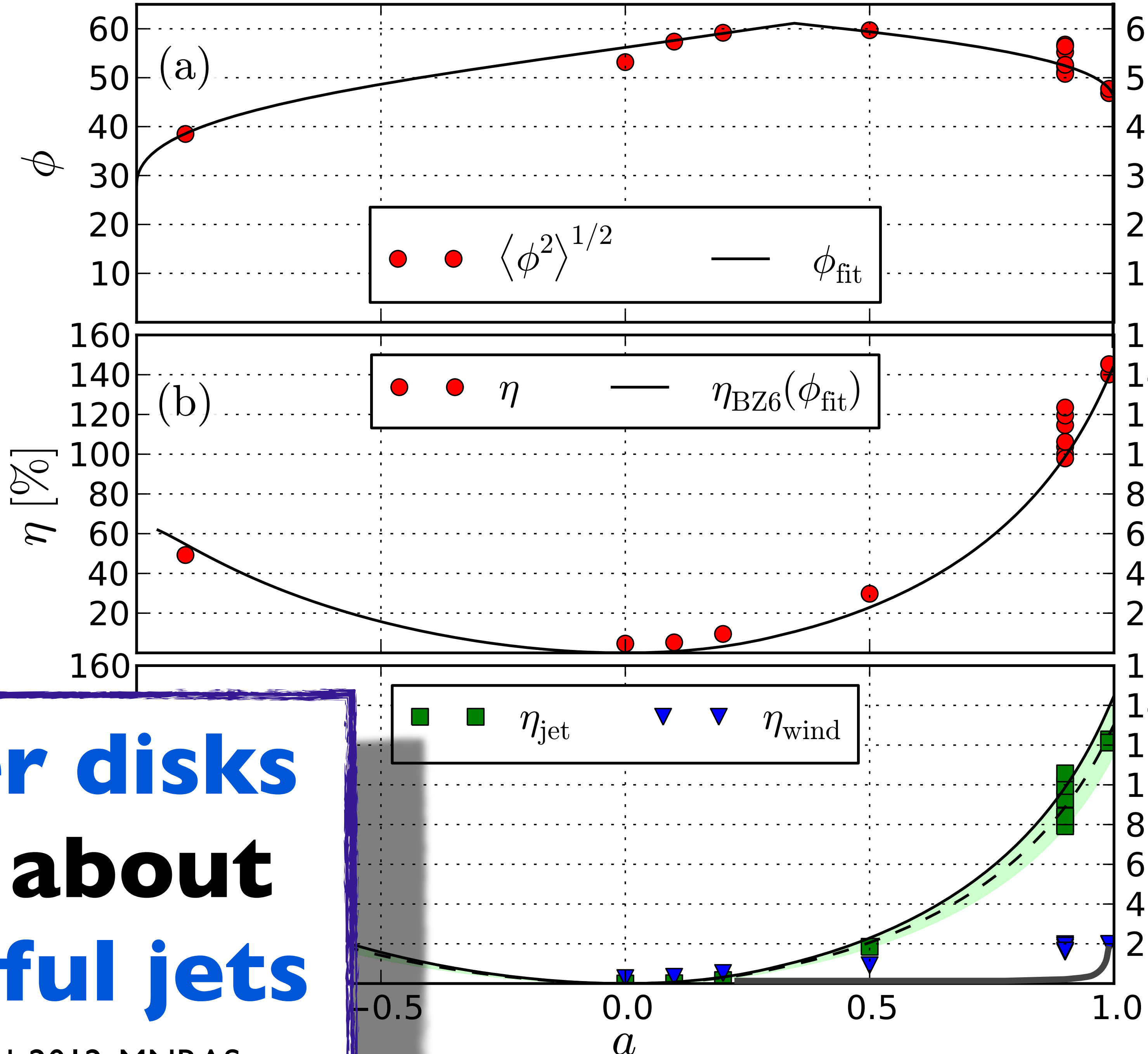

2x thicker disks produce about 2x powerful jets

(McKinney, AT, Blandford, 2012, MNRAS, submitted, arXiv:1201.4163)

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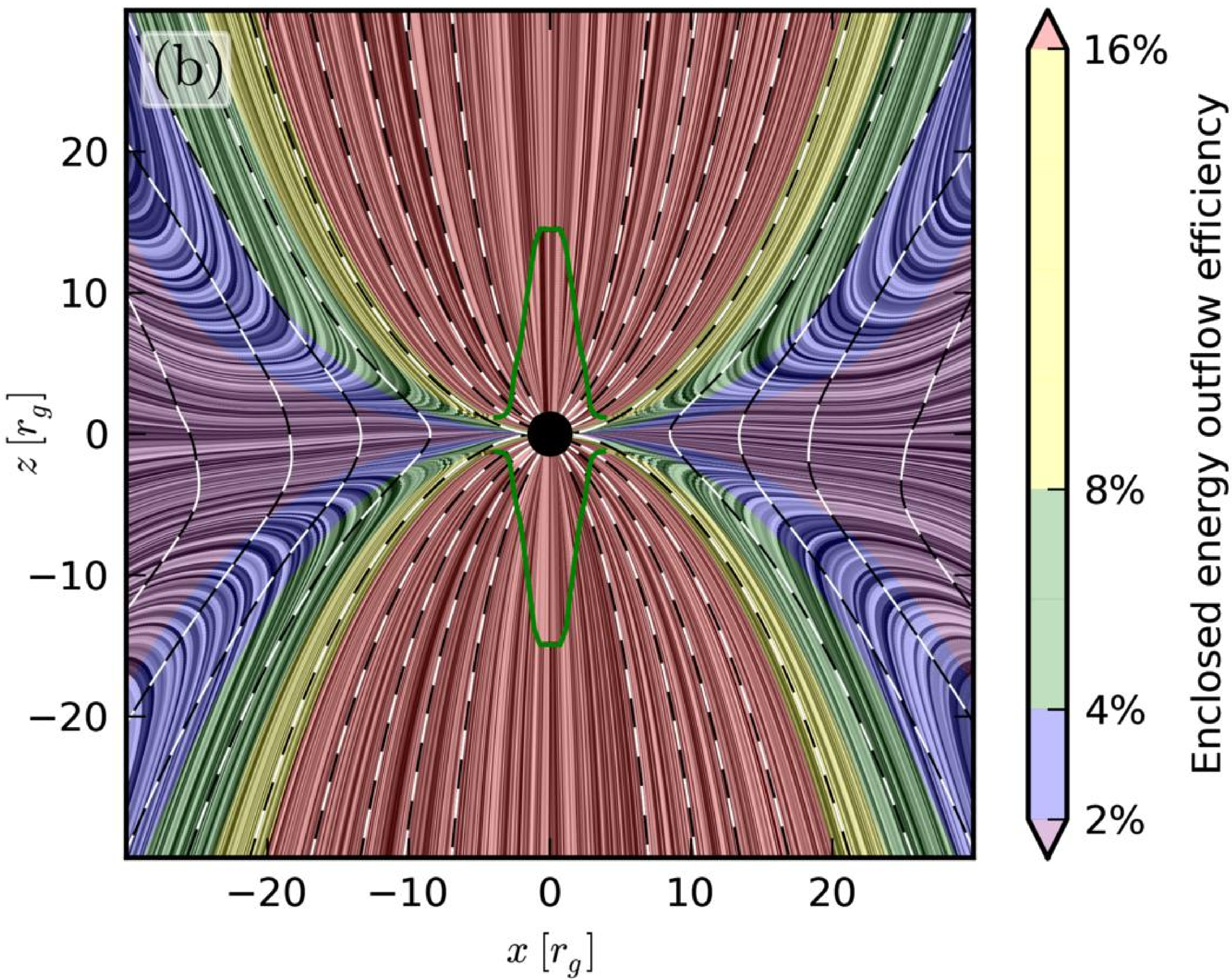
Where energy comes from?

$a = 0.9,\ \eta \approx 102\%,\ \eta_{\rm wind} \approx 16\%$



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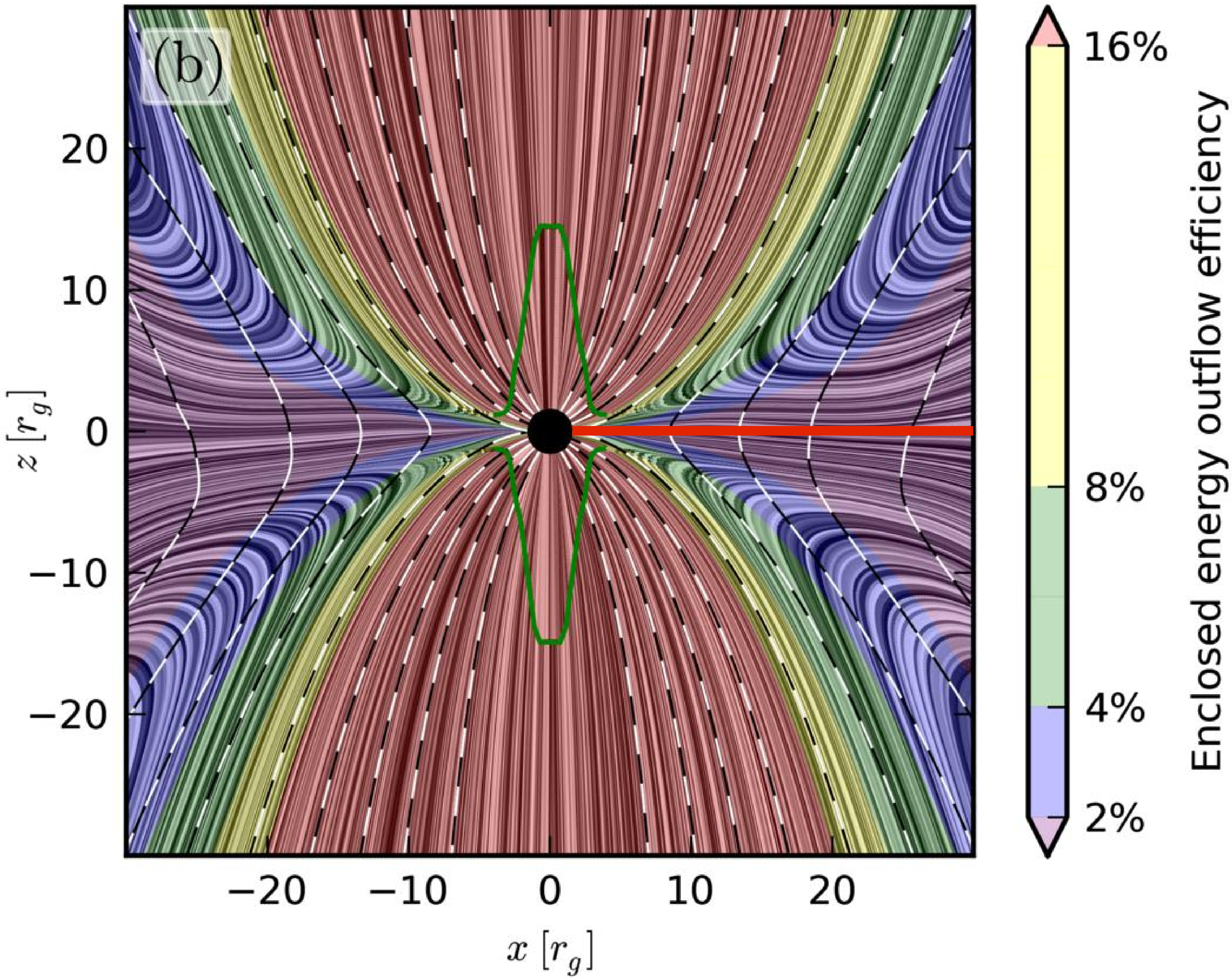

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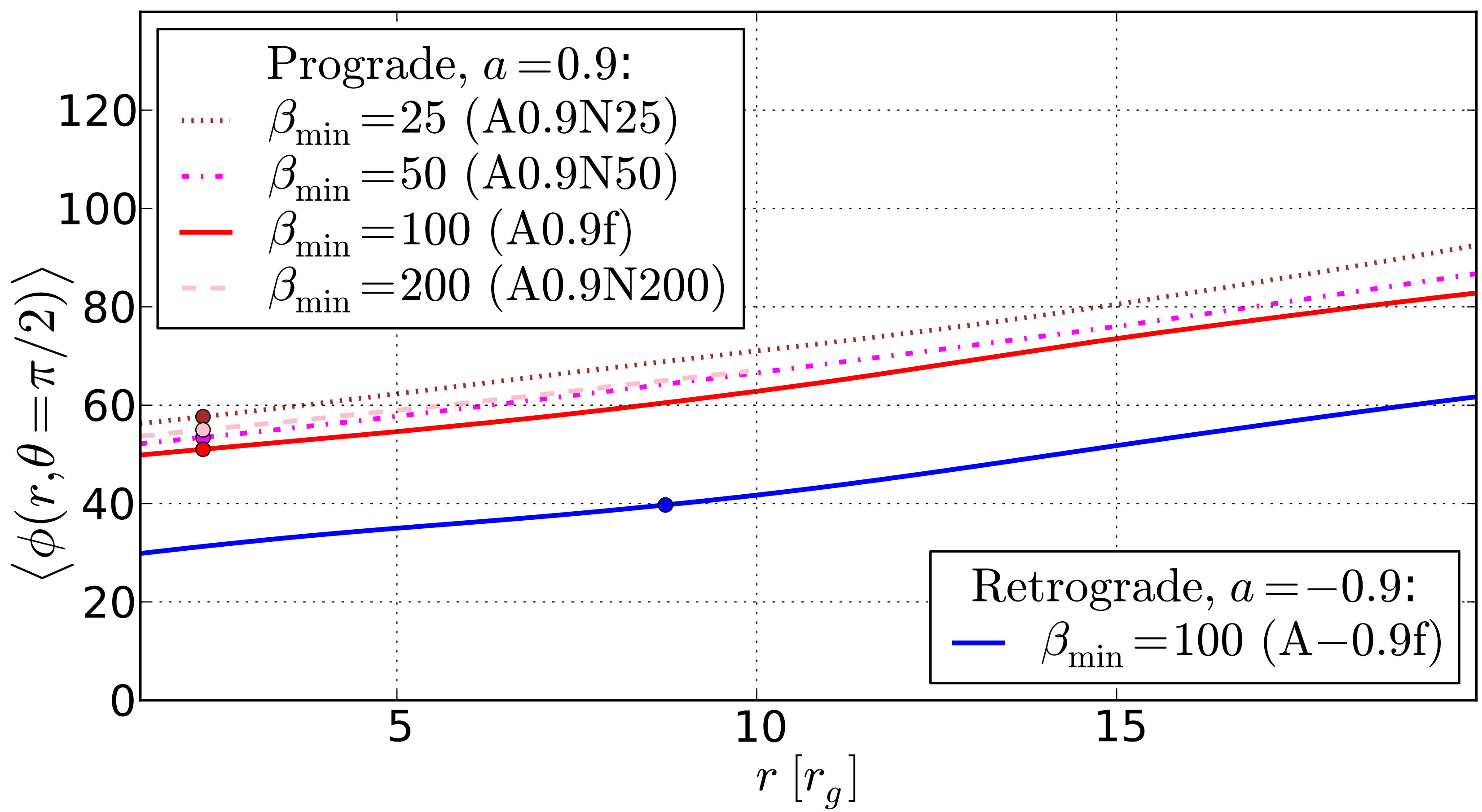
Initial flux does not matter!

Variation in field strength by 300% leads to variation less than 10% in BH flux and 20% in η



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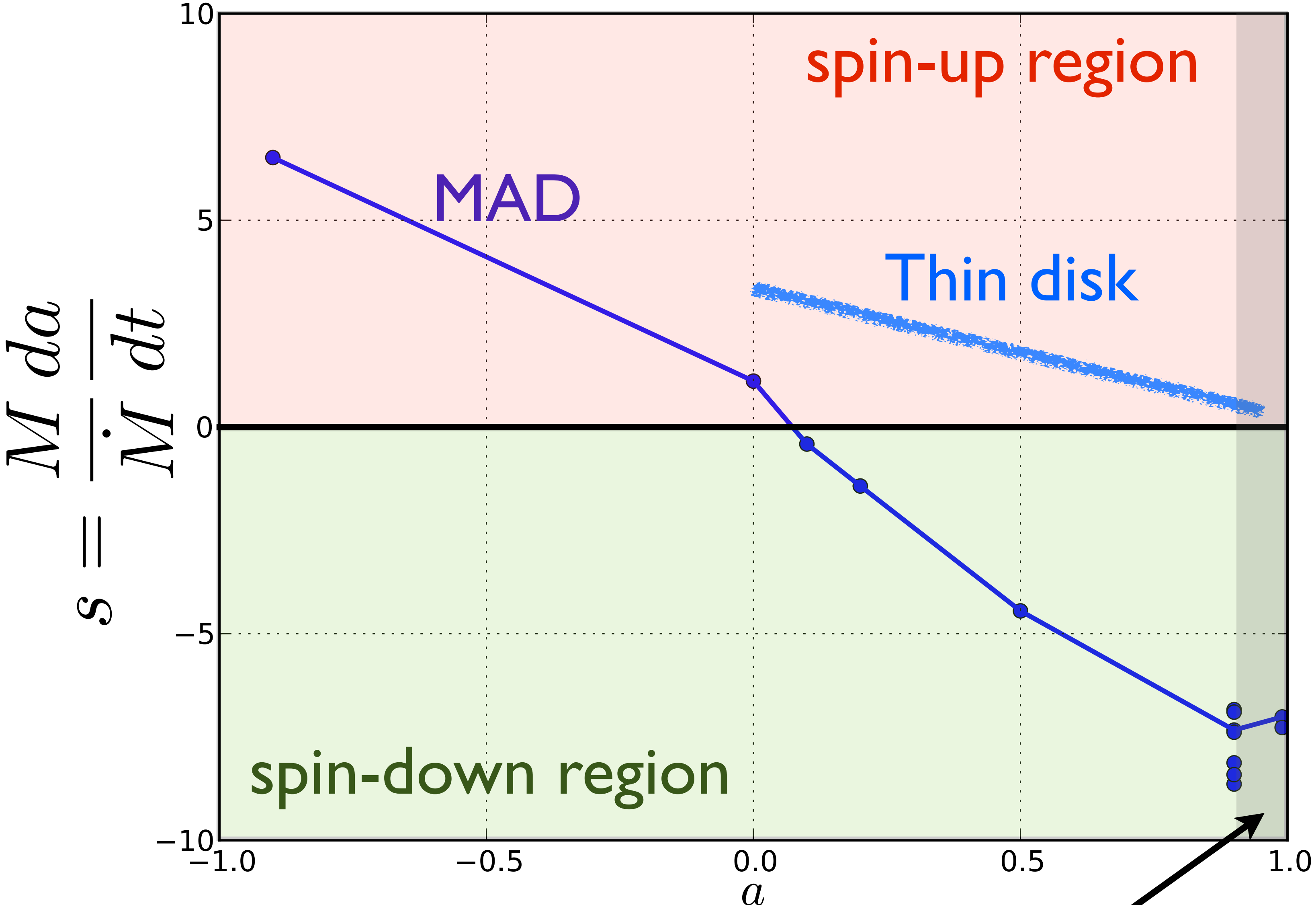
Our MADs slow BHs down to a halt



Conventional spin equilibrium region, $a \gtrsim 0.9$

(AT, McKinney 2012a, MNRAS, accepted, arXiv:1201.4385, 2012b, in prep.)

(see also Gammie et al. 2005, Shapiro et al. 2005, Benson & Babul 2009)

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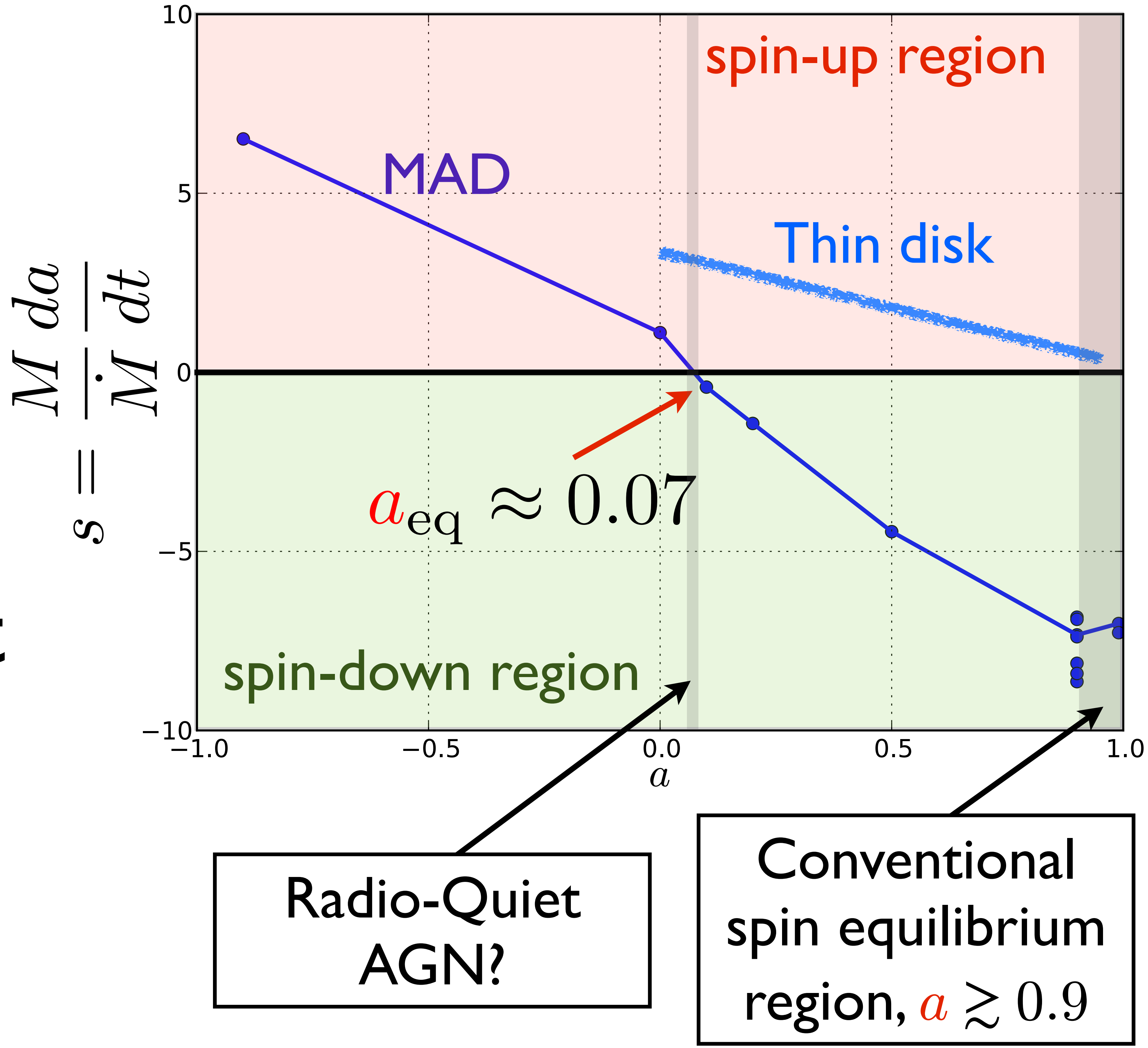

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Our MADs slow BHs down to a halt

In 1/3 of black hole mass doubling time,

$$\tau \approx 5\times 10^8 \frac{0.1}{\lambda_{\rm Edd}} \ [\rm yr]$$

BHs are spun down from

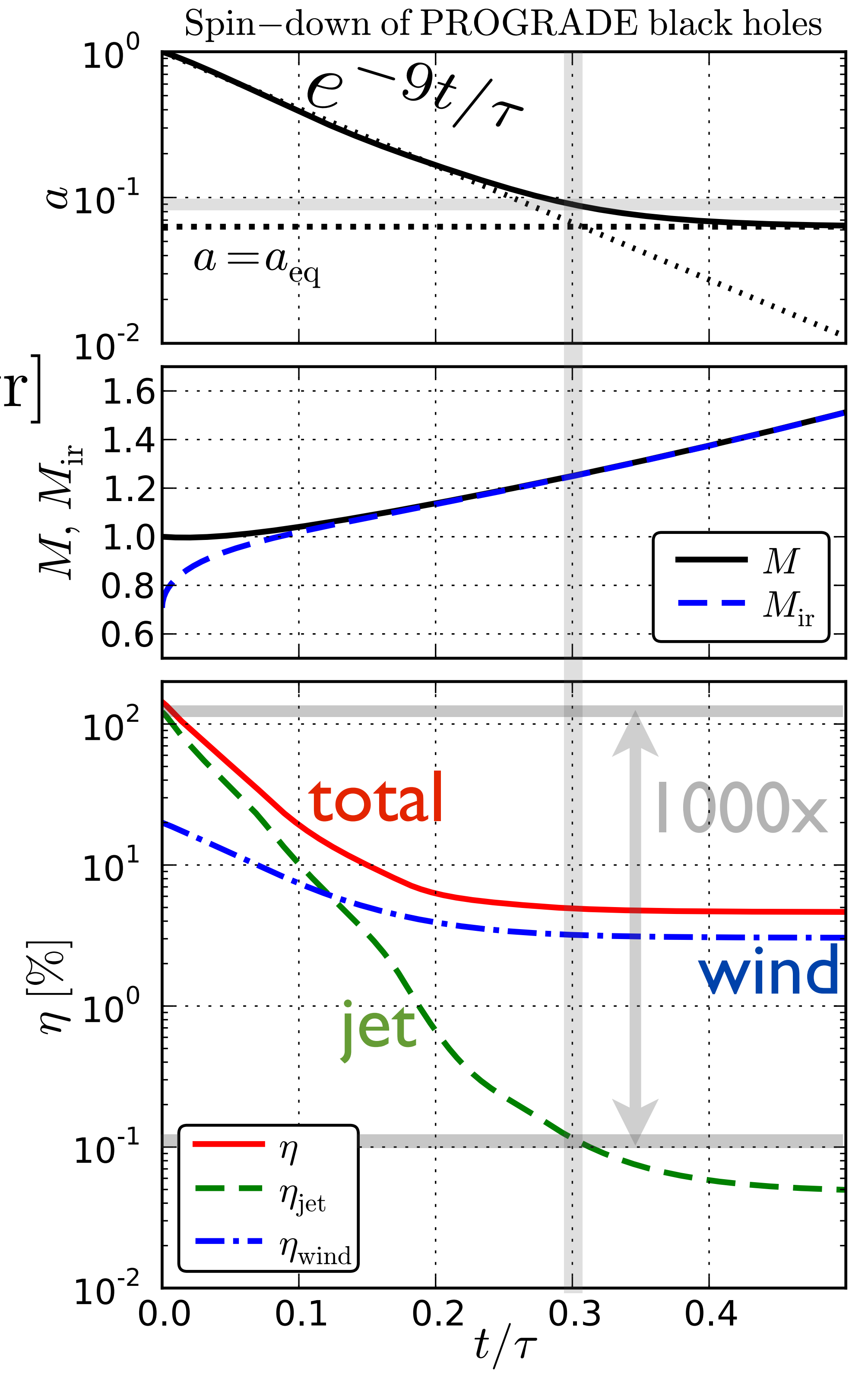
$$a = 1$$

to

$$a \lesssim 0.1.$$

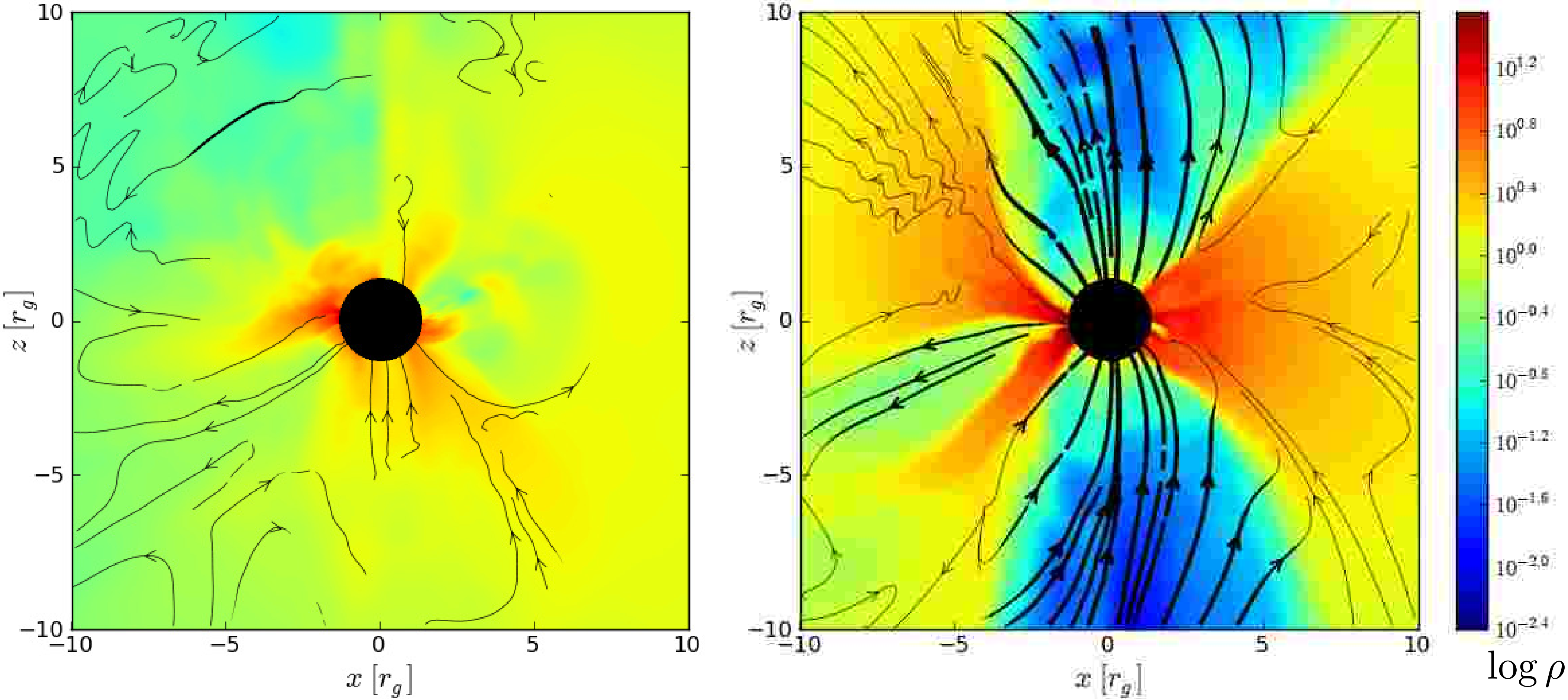
Jet efficiency drops by >1000x from >100% to 0.1%

(AT, McKinney 2012b, in prep.)



Anti-MADs: Zero Initial Flux

McKinney, AT, Blandford, 2012, MNRAS, in press, arXiv:1201.4163



$t \approx 40164 r_g/c$ $t \approx 62828 r_g/c$

$$\eta \approx 0.01\%$$

Is This The Radio-Quiet AGN?

MAD Summary

- Simulations maximize $\Phi \rightarrow$ MAD state:
  - $\eta > 100\%$ for $a \gtrsim 0.9$
  - MADs slow BHs down to a "halt", $a \lesssim 0.1$
- Radio-Loud AGN: MADs with $a \approx 1$
- Radio-Quiet AGN:
  - MADs with "halted" BHs, $a \lesssim 0.1$, or
  - Low vertical flux accretion for any a
- Retrograde BHs $\rightarrow$2-3x less powerful jets
- Thicker disks $\rightarrow$ more powerful jets